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**Input Optics Subsystem
Design Requirements Document**

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Abstract

The design requirements for the Advanced LIGO Input Optics Subsystem include the requirements for the RF modulation of the light, acquisition and operation of the mode cleaner, mode matching of the light to the interferometer, and beam steering into the interferometer. The scope of the IO includes the following hardware: phase modulation Pockels cells, photodetectors and related protective shutter (if needed), cabling, angular actuators to maintain alignment into the mode cleaner, mode cleaner optics, fabrication of suspensions, Faraday isolator, and mode matching telescopes.

1 Introduction

1.1 Purpose

The purpose of this document is to describe the design requirements for the Input Optics (IO) subsystem. Primary requirements are derived ('flowed down') from LIGO principal science requirements. Secondary requirements, which govern Detector performance through interactions between IO and other Detector subsystems, have been allocated by Detector Systems Engineering.

1.2 Applicable Documents

1.2.1 LIGO Documents

- [1] LIGO-T010075, "Advanced LIGO Systems Design", P. Fritschel, et al.
- [2] LIGO-M990288, "LIGO 2 Conceptual Project Book"
- [3] LIGO-E990303, "Seismic Isolation Subsystem Design Requirements Document", P. Fritschel, et al.
- [4] LIGO-T010076, "Optical Layout for Advanced LIGO", D. Coyne
- [5] LIGO-T0900142-v1, "Pointing Requirements in Advanced LIGO", G. Mueller,
- [6] LIGO-T000053-01-D "Cavity Optics Suspension Subsystem Design Requirements Document, P. Willems, et al.

- [7] LIGO-E010613, Generic Requirements & Standards for Detector Subsystems
- [8] Upgrading the Input Optics for High Power Operation, LIGO- T060267-00-D
- [9] Modulators and Isolators for Advanced LIGO, LIGO-G060361-00-D

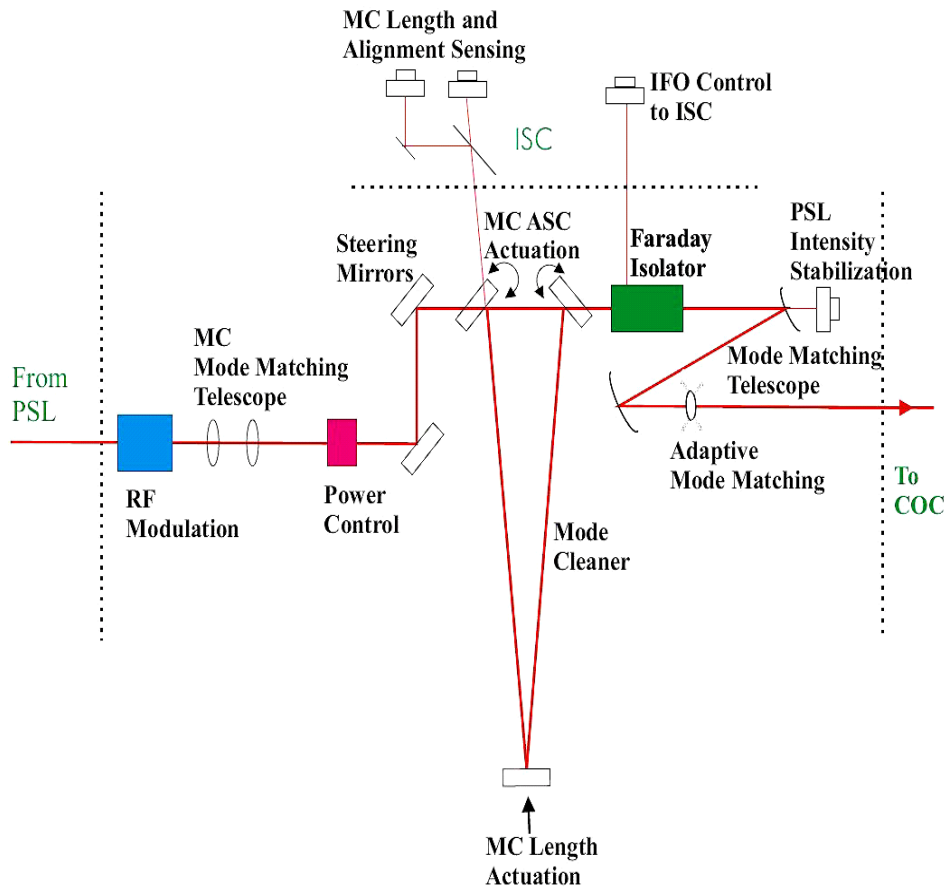
- [10] Advanced LIGO Input Optics Conceptual Design Document, LIGO T-020027-00-D
- [11] Advanced LIGO Input Optics Preliminary Design Document, LIGO T-060629-00-D
- [12] LIGO T050036-04-D, Pre-Stabilized Laser Design Requirements, B. Willke, et al.
- [13] LIGO T080195-01-D, Pre-Stabilized Laser Preliminary Design, P. King, et al.
- [14] LIGO-T070236-00-D, "Interferometer Sensing and Control Design Requirements, R. Adhikari et al.
- [15] LIGO T070247-01-I, AdvLIGO Interferometer Sensing and Control Conceptual Design, R. Abbott, et al.

2 General description

2.1 Product Functions

The IO conditions the laser light so that its properties are compatible with the primary scientific requirements for Advanced LIGO. This function separates into the following seven categories.

Figure 1 Overall IO Product Specification



2.1.1 RF Modulation

The ISC subsystems require frequency components of the laser light (sidebands) for controlling the interferometer. The IO must provide for the production and monitoring of these sidebands using RF signals from the ISC. Specifically, the IO provides:

- the design, fabrication, and assembly of all RF electro-optic modulators for the production of sidebands for the interferometer and mode cleaner length and alignment sensing systems, including ancillary optical and mechanical components as needed.

The IO does not provide:

- RF signals to drive the oscillators.

2.1.2 Mode Cleaning

The laser light must be frequency stabilized and spatially filtered before it can be used to provide length and alignment sensing for the IFO. This is accomplished through a mode cleaner cavity.

Specifically, the IO provides:

- the design, fabrication, and assembly of a suspended, in-vacuum triangular mode cleaner cavity for active frequency suppression through feedback to the PSL, passive frequency noise suppression above its cavity pole frequency, and passive spatial filtering at all frequencies.
- diagnostic beams needed for locking and aligning the mode cleaner.
- the design, fabrication, and assembly of ancillary optical and mechanical components for mode-matching and steering beams into and out of the mode cleaner.

The IO does not provide:

- the design and implementation of the mode cleaner length and alignment controls.
- the design of the mode cleaner and steering mirror suspensions.

2.1.3 Power Control

The interferometer will be operated at variable power levels during installation, commissioning, and operation. An actuator to adjust the optical power is required.

Specifically, the IO will provide:

- the design, fabrication, and assembly of a power adjustment system including the optical and mechanical components.

The IO does not provide:

- control signals for power adjustment
- intensity stabilization of the PSL

2.1.4 IFO Mode Matching

The light must be delivered to the IFO with a proper Gaussian mode so that it will resonate in the IFO and not be rejected. Thus the IO provides for the mode matching of the light between the mode cleaner and the core optics components of the interferometer. Specifically, the IO provides:

- the design, fabrication, and assembly of the pre-mode-matching telescope (PMMT) capable of accommodating small deviations from design specification in the core optics and providing optimal mode-matching for the different operational powers (currently 25 W and 125 W¹).

The IO does not provide:

- the design of the PMMT suspensions

2.1.5 Optical Isolation of the PSL and Distribution of ISC Signals

The IO must provide optical isolation between the COC and the PSL. In addition, the ISC requires optical pick-off signals for controlling the length and alignment of the IFO. Specifically, the IO provides:

- the design, fabrication, and assembly of an in-vacuum Faraday isolator capable of functioning over the entire power range of operation, including ancillary optical and mechanical components.
- diagnostic beams for other subsystems, including
 - a portion of the transmitted mode cleaner light for intensity stabilization of the PSL.
 - the reflected light from the power recycling mirror.

2.1.6 Internal Diagnostics

The IO must provide diagnostic capabilities for its own functions and for that of other subsystems where appropriate.

2.2 General Constraints

The IO is subject to the following constraints:

- The in-vacuum components of the IO must be contained in the current LLO and LHO vacuum envelopes (subject to the replacement of the HAM1,2 and 7,8 beam tube with a LIGO beam tube) and within the footprints of the Advanced LIGO seismic platforms.²

¹ LIGO-T070236-00-D, “Interferometer Sensing and Control Design Requirements, R. Adhikari et al.

² LIGO-T010076, “Optical Layout for Advanced LIGO”, D. Coyne

2.3 Assumptions and Dependencies

The following factors have been assumed in this document, and are, to the best of our knowledge, consistent with or have been flowed down from the Advanced LIGO Systems Design³ and LIGO II Conceptual Project Book⁴; (configuration control established TBD SYS).

2.3.1 Suspensions

- Suspensions provided by SUS
 - o Triple Pendula suspensions for MC mirrors
 - o LIGO 1 SOS for SM and PMMT mirrors

³ LIGO-T010075, “Advanced LIGO Systems Design”, P. Fritschel, et al.

⁴ LIGO-M990288, “LIGO 2 Conceptual Project Book”

3 Requirements

3.1 Introduction

The IO subsystem derives its requirements, where available, from the top-level Advanced LIGO requirements for sensitivity and availability. The requirements are grouped into sections corresponding to the following functions of the IOO: RF modulation, mode cleaner, optical isolation, power control, and IFO mode matching.

The accompanying conceptual design document will present corresponding design for meeting these requirements.

We derive all requirements below assuming that the related noise amplitude spectral density is held to 10% of the envelope of the LIGO strain sensitivities $h(f)$ at all in-band frequencies for all operational modes.

3.2 General IO Requirements

3.2.1 Optical Efficiency

The net efficiency of IO TEM₀₀ optical power transmission from PSL output to COC input shall be 0.75 or greater, determined by the requirement that at least 125 W of TEM₀₀ light be coupled into the IFO and assuming the PSL requirement of delivering > 165 W in a TEM₀₀ mode to the IO. The optical power of the IO is defined as the sum for the carrier and sidebands.

3.2.2 Output Beam In-band Alignment Stability (Jitter)⁵

Alignment fluctuations at the input of the COC couple to angular motion of the test masses to give in-band displacement signals. The (in-band) alignment stability of the entire IO subsystem shall not compromise that achieved directly after the mode cleaner, including the mode-matching telescope. The output beam alignment stability requirement is:

- (DC readout) Contributions from the carrier jitter coupled to ITM/ETM differential tilts dominate the noise contribution. We find:

$$\varepsilon_1(f) < 1.0 \times 10^{-8} \sqrt{1 + \left(\frac{100 \text{ Hz}}{f}\right)^4} \frac{[2 \times 10^{-9}]}{\Delta\Theta_{ITM}} \frac{1}{\sqrt{\text{Hz}}}$$

where ε_j is the amplitude of the 10,01 modes, and $\Delta\Theta_{ITM} = (\Theta_{ITM1} - \Theta_{ITM2})$ is the total differential tilt of the ITMs.

3.2.3 Parasitic Interferometers

Light that reflects or scatters into the Rayleigh angle of the beam contributes to the in-band signal directly (through phase modulation at GW frequencies) and indirectly (through frequency shifting of scattered and reflected light into GW band due to mirror

⁵ LIGO-T0900142, "Pointing Requirements in Advanced LIGO", G. Mueller,

motions). Carrier light which propagates through transmissive optics can be upconverted to in-band noise through seismic excitation via the Fizeau effect. Carrier and sideband light reflecting off of mirrors can be directly upconverted to in-band signals through the Doppler effect. In addition, spurious cavities can be formed by light that reflects back to normal incidence optical components.

We require:

- For optical components located on the PSL table, the frequency noise must be limited to 10% PSL frequency noise specification⁶:
 - o $\delta\nu(f) < 0.1 \text{ Hz/Hz}^{1/2}$ $f < 20 \text{ Hz}$
 - o $\delta\nu(f) < 0.1 (20\text{Hz}/f) \text{ Hz/Hz}^{1/2}$ $20 \text{ Hz} < f < 1 \text{ kHz}$
 - o $\delta\nu(f) < 0.002 \text{ Hz/Hz}^{1/2}$ $f > 1 \text{ kHz}$
- For optical components located after the MC, the frequency noise must be limited to 10% of the MC requirement (see 3.4.1):
 - o $\delta\nu(f) < 5 \times 10^{-3} \text{ Hz/Hz}^{1/2}$ at 10 Hz
 - o $\delta\nu(f) < 1 \times 10^{-4} \text{ Hz/Hz}^{1/2}$ at 100 Hz

3.3 RF Modulation Requirements

The IO provides the RF modulation for the optical sidebands used in the length and alignment sensing. The requirements include modulation frequencies, modulation depths, and relative stability of the mode cleaner resonance and modulation frequency and amplitude.

3.3.1 Modulation Frequencies

The IFO requires frequencies that resonate in the interferometer and mode cleaner as well as for alignment sensing and control. All frequencies (chosen by ISC) must pass through the mode cleaner and therefore be integral multiples of the mode cleaner free spectral range.

- The IO must provide RF modulation at frequencies consistent with the Advanced LIGO Sensing and Control scheme. The current set of frequencies span a range from 8.5 MHz to 45.5 MHz.

3.3.2 Modulation Depths

Modulation depths depends on the final length and alignment sensing and control scheme. Based on the conceptual design, the following range is sufficient:

- The IO must provide for modulation depths in the range $m = 0-0.8$.

⁶ LIGO-T070247-01-I “AdvLIGO Interferometer Sensing and Control Conceptual Design, R.Abbott et al.

3.3.3 Modulation Amplitude Noise Requirements⁷

Fluctuations in the amplitude of the modulation sidebands introduce noise in the servo system and also change the intensity of the carrier which then changes for example the radiation pressure noise in the arm cavities. We require:

- RAN SB (f) < $10^{-7} \text{ Hz}^{-1/2}$ $f < 10 \text{ Hz}$
- RAN SB (f) < $10^{-7} \text{ Hz}^{-1/2} (10\text{Hz}/f)$ $10 \text{ Hz} < f < 100\text{Hz}$
- RAN SB (f) < $10^{-8} \text{ Hz}^{-1/2} (100\text{Hz}/f)^{1/2}$ $100 \text{ Hz} < f < 1 \text{ kHz}$
- RAN SB (f) < $3 \times 10^{-9} \text{ Hz}^{-1/2}$ $f > 1 \text{ kHz}$

Note: These “requirements” are equal to the noise inputs used in LIGO-T070247-01-I. Based on the generated noise budgets, these values can be relaxed at higher frequencies. The requirements presented in LIGO-T070236-00-D⁸ are based on a different IFO design.

3.3.4 Modulation Phase Noise Requirements⁷

Fluctuations in the modulation frequency due to oscillator phase noise beat at the length sensing port photodetectors and produce technical noise.

- $\Phi_{\text{OSC}} < 10^{-5} \text{ rad/Hz}^{1/2} (10 \text{ Hz}/f)^{3/2}$ $10 \text{ Hz} < f < 100\text{Hz}$
- $\phi_{\text{OSC}} < (3 \times 10^{-7}) \text{ rad/Hz}^{1/2} (100 \text{ Hz}/f)$ $100 \text{ Hz} < f < 1 \text{ kHz}$
- $\phi_{\text{OSC}} < (3 \times 10^{-8}) \text{ rad/Hz}^{1/2}$ $f > 1 \text{ kHz}$

Note: These “requirements” are equal to the noise inputs used in LIGO-T070247-01-I. Based on the generated noise budgets, it should be possible to relax these values.

3.3.5 Modulation Cross Products

There are three modulation frequencies in use (including the modulation for the mode cleaner. The modulation process produces intermodulation products (sidebands on sidebands). These intermodulation products can mix down into the GW band and produce technical noise. These modulation cross products look like fluctuations in the modulation frequency and have identical requirements:

- $\Phi_{\text{OSC}} < 10^{-5} \text{ rad/Hz}^{1/2} (10 \text{ Hz}/f)^{3/2}$ $10 \text{ Hz} < f < 100\text{Hz}$
- $\phi_{\text{OSC}} < (3 \times 10^{-7}) \text{ rad/Hz}^{1/2} (100 \text{ Hz}/f)$ $100 \text{ Hz} < f < 1 \text{ kHz}$
- $\phi_{\text{OSC}} < (3 \times 10^{-8}) \text{ rad/Hz}^{1/2}$ $f > 1 \text{ kHz}$

3.4 Mode Cleaner Requirements

The mode cleaner provides frequency and spatial stabilization of the laser light, as well as intensity stabilization above its pole frequency.

⁷ LIGO-T070247-01-I, “AdvLIGO Interferometer Sensing and Control Conceptual Design”, R. Abbott et al.

⁸ LIGO-T070236-00-D, “Interferometer Sensing and Control Design Requirements, R. Adhikari et al.

3.4.1 Mode Cleaner Frequency and Intensity Stabilization

The frequency noise requirements of the stabilized laser beam⁹ together with the common mode servo gain set the frequency noise requirements for the mode cleaner. The frequency noise requirement of the stabilized beam at 100 Hz is $1 \times 10^{-6} \text{ Hz/Hz}^{1/2}$, the iLIGO common mode servo gain at 100 Hz is 10^4 leading to a frequency noise requirement of the interferometer input beam of $1 \times 10^{-2} \text{ Hz/rtHz}$. An additional safety factor of 10 is used to allow to reduce the common mode servo gain.

We require:

- $\delta v(f) < 5 \times 10^{-2} \text{ Hz/Hz}^{1/2}$ at 10 Hz
- $\delta v(f) < 1 \times 10^{-3} \text{ Hz/Hz}^{1/2}$ at 100 Hz

This limits the changes in the mode cleaner length (Half Trip) to be below:

- $\delta L(f) < 2.5 \times 10^{-15} \text{ m/Hz}^{1/2}$ at 10 Hz
- $\delta L(f) < 5 \times 10^{-17} \text{ m/Hz}^{1/2}$ at 100 Hz

The mode cleaner length changes (half trip) caused by technical radiation pressure noise are:

- $\delta L(f) = 5.7 \times 10^{-17} \text{ m/Hz}^{1/2} (\text{RIN}(f)/[2 \times 10^{-9}/\text{Hz}^{1/2}]) ([10\text{Hz}]/f)^2$

Note that the relative intensity noise used in this equation is white (shot noise limited) and assumed to be equal to the RIN going into the main interferometer. The later is unlikely the case, however we do not have a good model for the power fluctuations induced by the IO components between the MC and the PRM. Nevertheless, the margin of 2 orders of magnitude should be sufficient to ensure that the additional RIN inside the MC is not compromising the frequency stability of the MC.

3.4.2 Mode Cleaner Alignment

The MC alignment requirement is given by:

- Low frequency – the frequency stability of the laser field must not be compromised by beam jitter, giving:

$$\Delta\Theta_{\text{MC}} < 3 \times 10^{-5} \text{ rad(rms)}$$

Note: This value depends on the Common mode gain. It is consistent with the requirements at 100Hz (see 3.4.1 for details). A detailed analysis would require a reliable spectrum for beam jitter and the frequency dependency of the common mode servo.

- In-band. There are two effects which limit the allowed angular motion of the MC mirrors:

⁹ “Interferometer Sensing and Control Design Requirements”, LIGO-T070236-00-D, R. Adhikari, et al.

1. The angular motion induces a length change if the laser beam is not centered on the rotational axis (piston effect). Assuming a beam centering of 1 mm (see 3.4.4), we require:

$$\delta\Theta_{MC}(f) < 2 \times 10^{-14} \sqrt{1 + \left(\frac{100 \text{ Hz}}{f}\right)^4} \frac{[2 \times 10^{-9} \text{ rad}]}{\Delta\Theta_{ITM}} \frac{\text{rad}}{\sqrt{\text{Hz}}}$$

2. The MC jitter rejection must not be compromised by MC mirror angular fluctuation noise (see 3.2.2). The induced beam jitter scales with the inverse beam size on the MC mirrors. This leads to requirements which are ~ 2 orders of magnitude less stringent.

3.4.3 Mode Cleaner FSR stability

Changes in the FSR require a readjustment of the demodulation phases to zero offsets in the error signals. The effect on the signal cannot be evaluated without a detailed length sensing and control scheme (TBD).

3.4.4 Mode Cleaner Beam Centering

The beam spot must be centered in the mode cleaner mirrors to a precision of 1 mm to avoid length-misalignment couplings.¹⁰

3.4.5 Availability Requirements

The IO availability will be limited by the lock acquisition time of the mode cleaner, and any degradation in performance due to thermal stress or optical contamination. We require:

- **Lock acquisition time to fully operational state < 20 sec**

3.5 Optical Isolation Requirements

Optical isolation is required to separate the MC from the IFO back reflected light. Based on evidence for parametric interferometers in initial LIGO, we can infer a requirement by power scaling arguments. Assuming an input of 125 W to the FI and a mode cleaner finesse of 500, and allowing a safety factor:

- **The required isolation level is 30 dB.**

3.6 Power Control

The Input Optics shall provide adjustable power to the interferometer for diagnostic and operational functions. We require:

- power range: 0 – 165 W (full PSL power)

¹⁰ LIGO-T000053-01-D “Cavity Optics Suspension Subsystem Design Requirements Document, P. Willems, et al.

- that the rate of the power change (dP/dt) be sufficiently small so as to not break the MC or IFO lock via radiation pressure kicks (both length and alignment)
- that the power control optics do not compromise the frequency or amplitude performance of the PSL.
- The power control system (hardware) will be designed with the flexibility to diagnostically vary the input power without substantial impairment of the interferometer's function, and without invasive alteration of the optics.

3.7 IFO Mode Matching

3.7.1 Mode Matching Coupling Efficiency

The coupling efficiency from the Input Optics to the Main Interferometer GW carrier and sidebands TEM₀₀, mode parameters as described in interfaces, (COC) shall be 0.95 or higher. The telescope will provide this level of coupling with adjustability to accommodate deviations in COC specifications and thermal lensing effects within the Input Optics.

3.7.2 Beam Steering

The IO MMT is used to steer the main beam into the IFO. For alignment control purposes, we require that the MMT design

- Preserve translation and angular adjustments into the IFO.

3.7.3 Mode Matching Telescope Alignment

Perturbations of the mode matching telescope may enhance the coupling of noise sources to gravitational wave noise (in band) and reduce coupling efficiency into COC (low frequency). We require that any telescopic magnification of pointing drift or jitter does not compromise the alignment stability of the IO output beam into the COC:

- Low frequency drift: telescope pointing must be consistent with COC coupling efficiency requirements of 0.95 or better:
 - the IO will furnish a diagnostic capability for measuring $\Delta\Theta_{MMT}$ for each MMT mirror at a level consistent with this requirement.
- Low frequency mode mismatch
 - the IO will furnish a diagnostic capability for determining Δw , the deviation from the ideal arm cavity beam waist size, and $\Delta\theta$, the deviation from the ideal arm beam divergence angle
 - the IOO will provide a method for in-situ correction of Δw and $\Delta\theta$
- In-band noise: telescope angular and displacement fluctuations must be consistent with ASC requirements for beam jitter at the input of the COC (see 3.2.2):

$$\delta\Theta_{MMT}(f) < 1 \times 10^{-12} \sqrt{1 + \left(\frac{100 \text{ Hz}}{f}\right)^4} \frac{[2 \times 10^{-9} \text{ rad}] \text{ rad}}{\Delta\Theta_{ITM} \sqrt{\text{Hz}}}$$

3.8 Baffles

The responsibility for the baffles in the IO region has recently been moved from AOS to IO. The AOS subsystem has released a conceptual design document (T060263) that includes these baffles, but IO will create its own scattered/stray light control document that presents these baffles in the context of the IO subsystem.

3.9 Optical Levers

In contrast to the original design, none of the IO-mirrors will be monitored with optical levers. Instead, the HAM tables in the IO area will be monitored by optical levers. The optical lever transmitters, receivers, and viewports are the responsibility of AOS.

3.10 Diagnostics

The diagnostic mode will provide the means to determine the proper functioning of the IO, and provide measurement of the performance of other subsystems. The following diagnostic capabilities are required of the IO (does not include diagnostics for MC ISC):

- IO Diagnostics
 - internal Qs of MC mirrors
 - Qs of MC suspensions
 - diagonalization state of MC optics
 - MC spatial stabilization
 - MC alignment and leveling
 - MC storage time
 - MC response to laser light pointing, frequency and intensity modulation
 - modulation sideband amplitudes and frequencies
 - RFAM level
 - MC beam centering
 - IFO mode matching efficiency
 - MC FSR stability
- Diagnostic Services
 - open loop mode cleaner mirror seismic excitation
 - variations in RF sideband modulation depth
 - variations in IFO mode matching efficiency
 - sideband detuning from mode cleaner resonance
 - variation in optical power

3.11 Other Requirements

The Input Optics shall meet all generic subsystem requirements as documented in Generic Requirements & Standards for Detector Subsystems, LIGO-E010613, D. Coyne.