GRAVITATIONAL WAVES: A NEW WINDOW ON THE UNIVERSE

Stephen Fairhurst

for the LIGO Scientific and Virgo Collaborations

References

Phys. Rev. X **6**, 041015 (2016)

GW150914: PRL 116, 061102 (2016)

GW151226: PRL 116, 241103 (2016)


Manhattan Mini Storage.com
212-storage
Scientists found gravitational waves in outer space.
If only it were that easy to find a walk-in closet in NYC.
Get your own Personal Closet:
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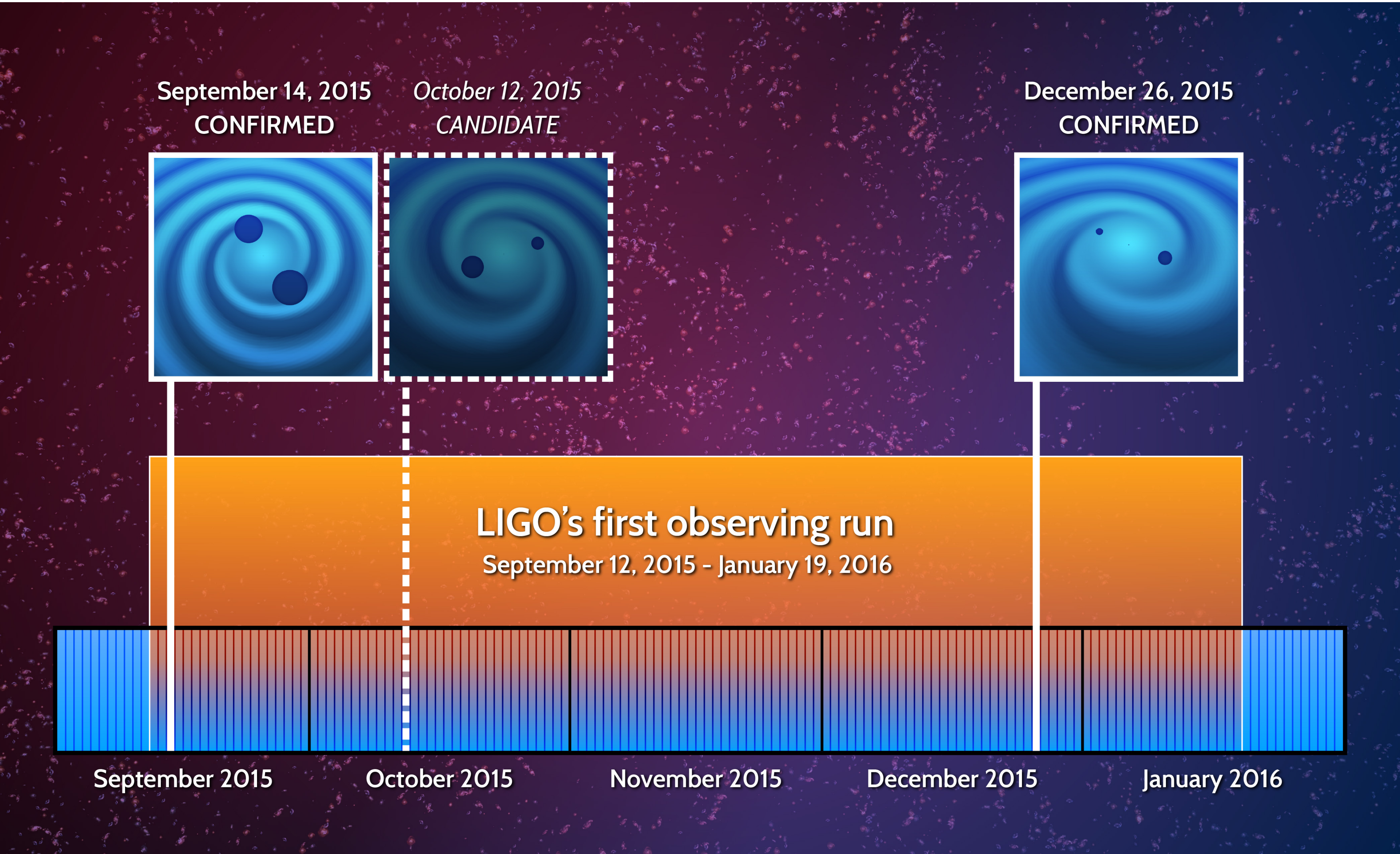
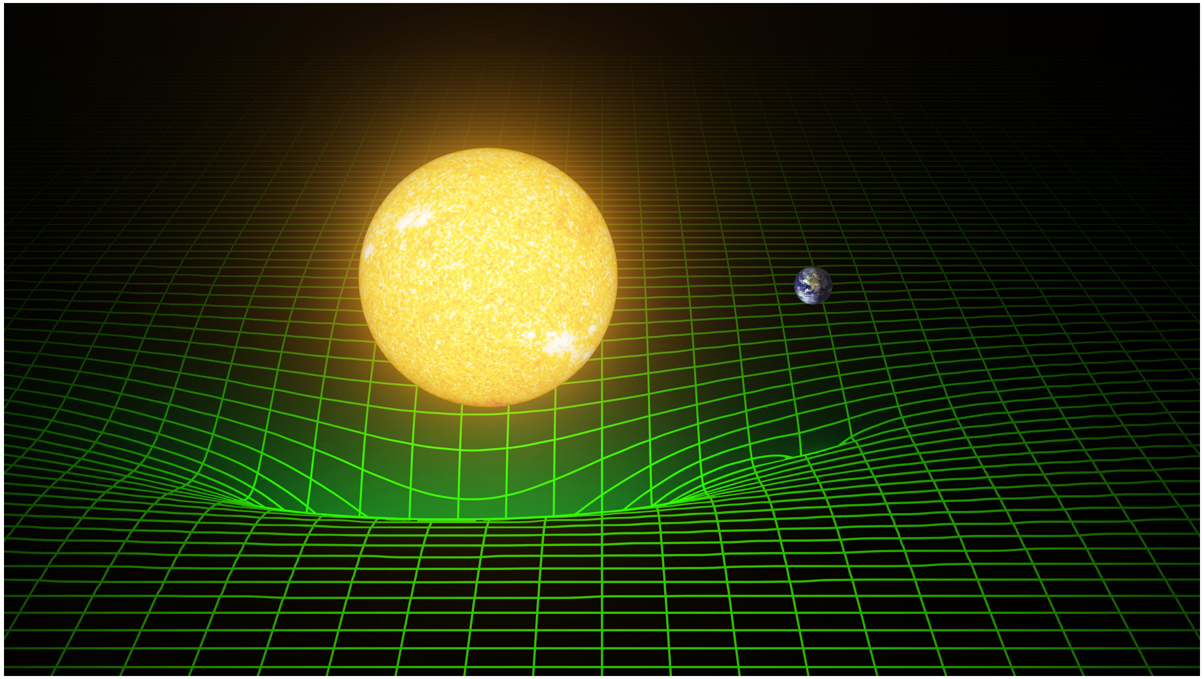
LIGO


September 14, 2015
CONFIRMED
October 12, 2015
CANDIDATE
December 26, 2015
CONFIRMED
LIGO's first observing run
September 12, 2015 - January 19, 2016
September 2015
October 2015
November 2015
December 2015
January 2016

General Relativity

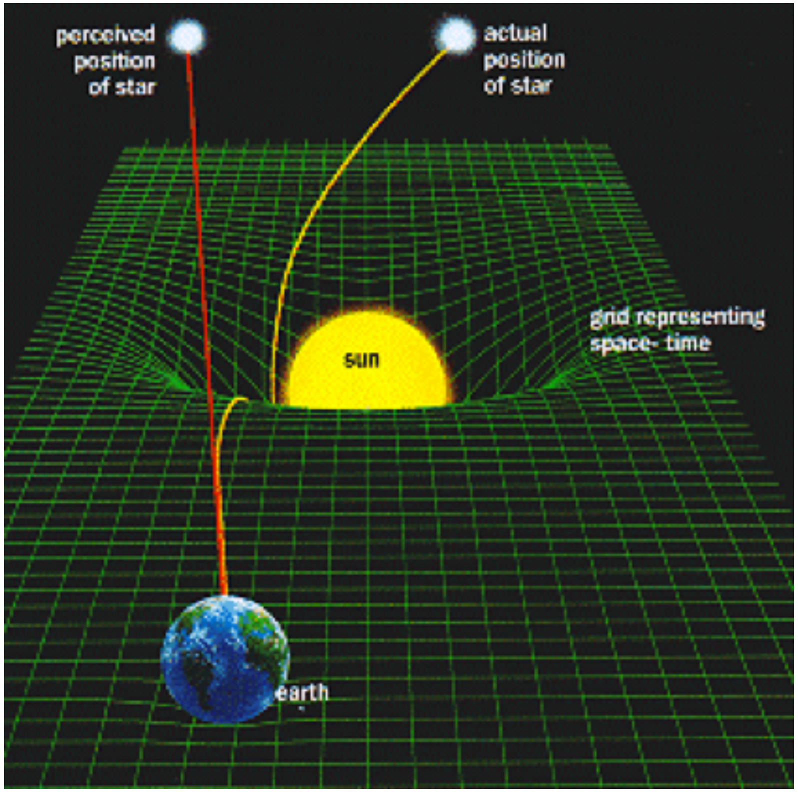
$$R_{\mu\nu} - \frac{1}{2}Rg_{\mu\nu} = 8\pi G T_{\mu\nu}$$

Matter tells space how to curve.
Space tells matter how to move.
- John Wheeler



General Relativity

$$\frac{D^2 X^\mu}{dt^2} = R^\mu{}_{\nu\rho\sigma} T^\nu T^\rho X^\sigma$$



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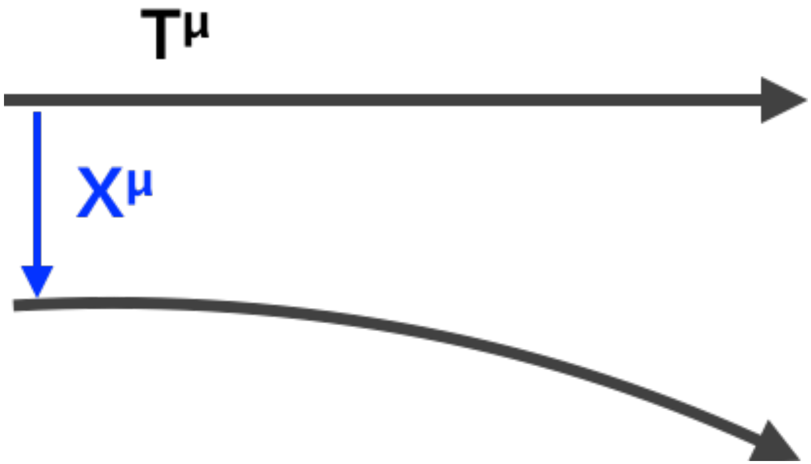
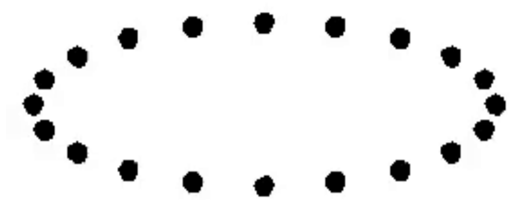


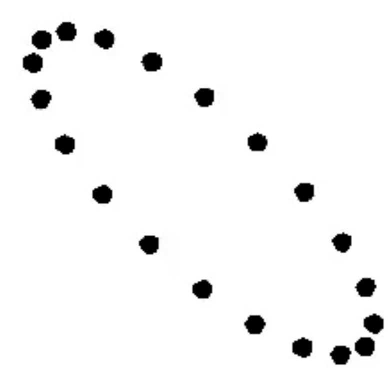
Linearized gravity

- Flat, empty space is a solution to general relativity.

- Leading order correction, $h_{\mu\nu}$, satisfies wave equation

- These waves create a tidal distortion in space-time, $h = \frac{\delta L}{L}$





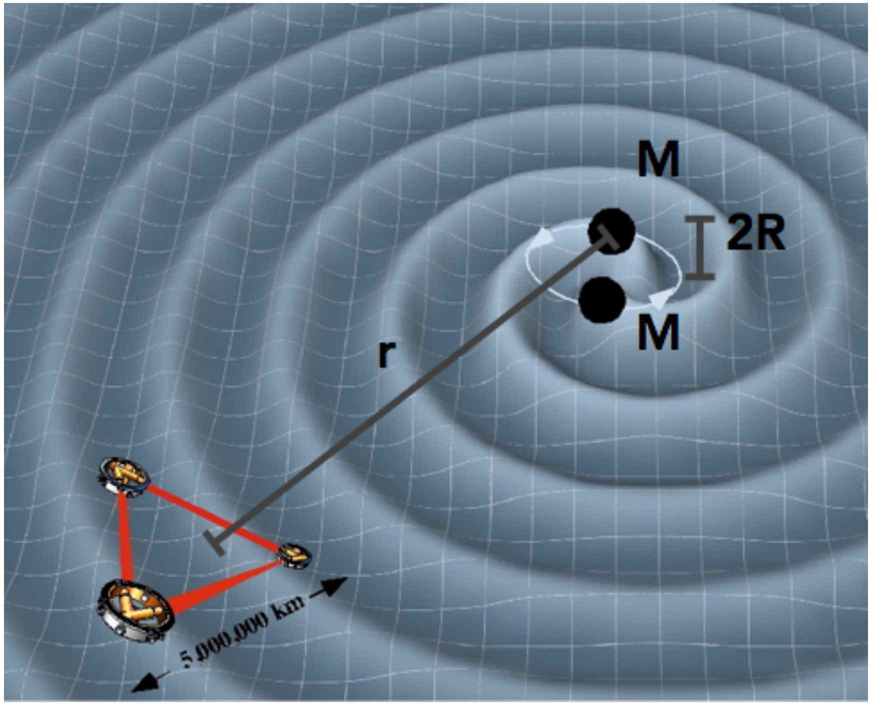
Generating gravitational waves

- Time varying mass quadrupole generates gravitational waves

- Binary system is ideal

$$h \sim \left(\frac{GM}{c^2 R}\right)\left(\frac{GM}{c^2 r}\right)$$

$$P \sim \frac{GM^2 v^6}{c^5 R^2}$$



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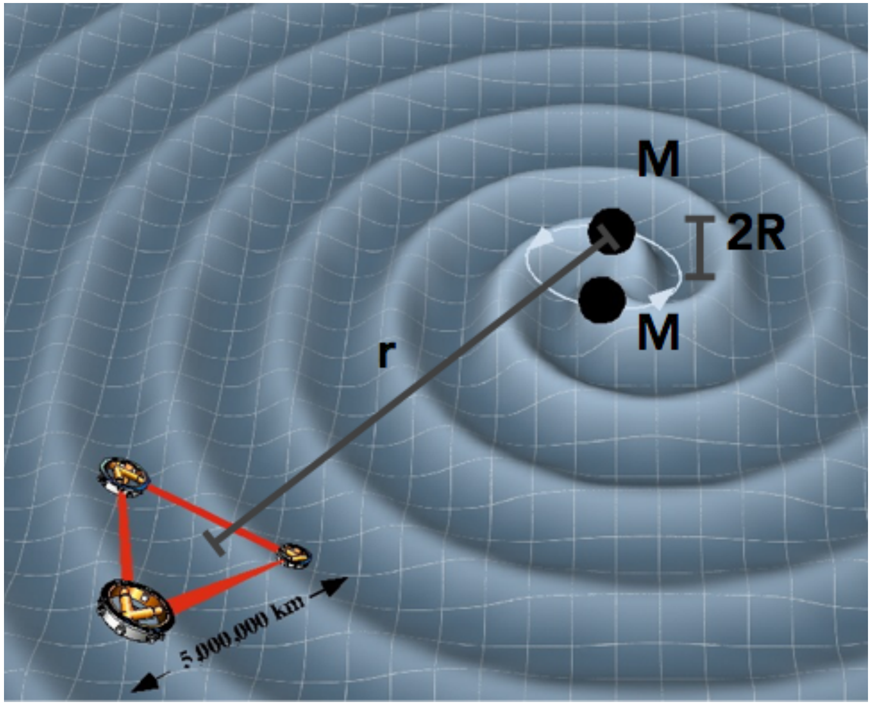
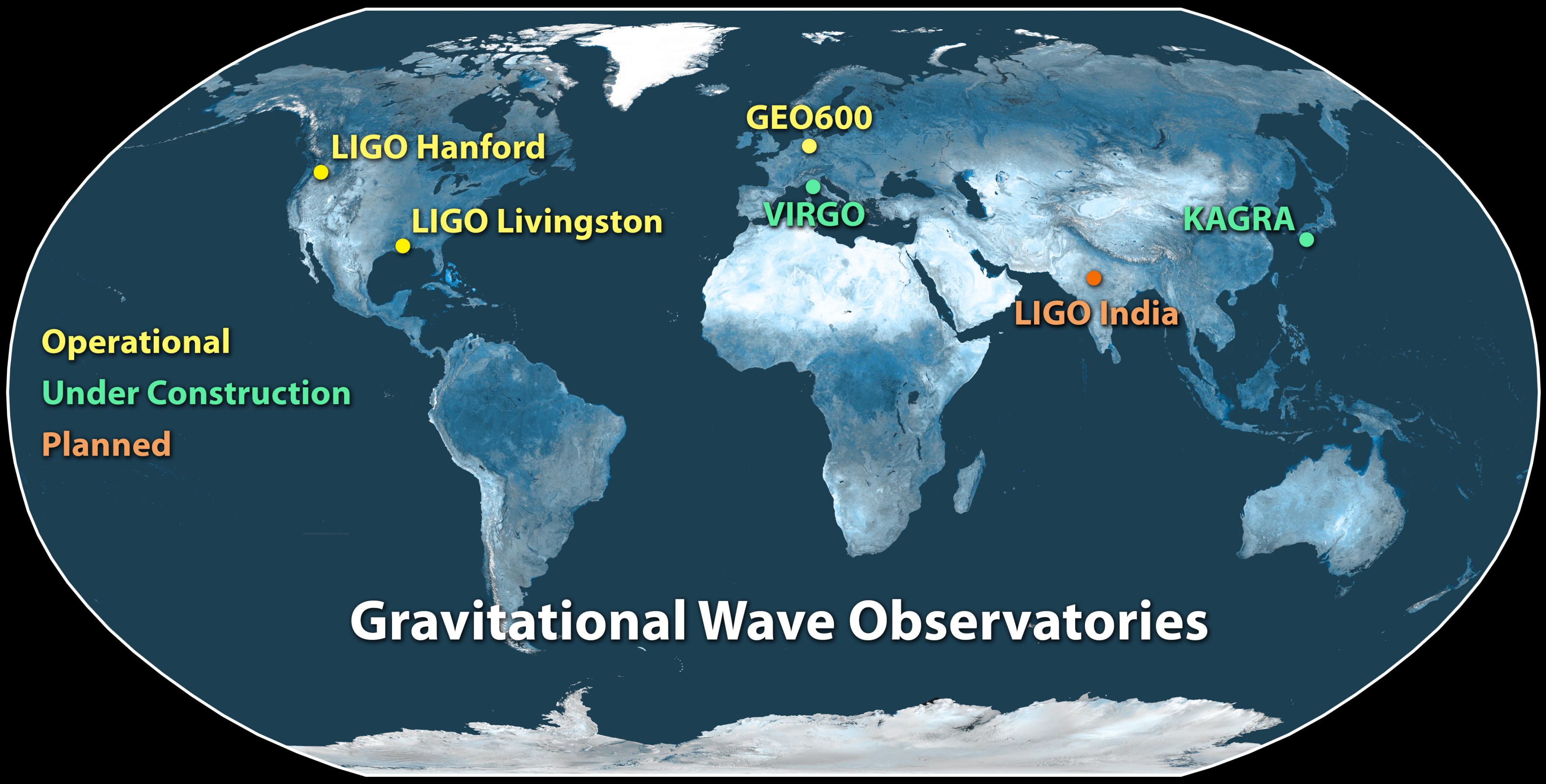


For a black hole:

$$R_{\rm Sch} = \frac{2GM}{c^2}$$

A global network



LIGO

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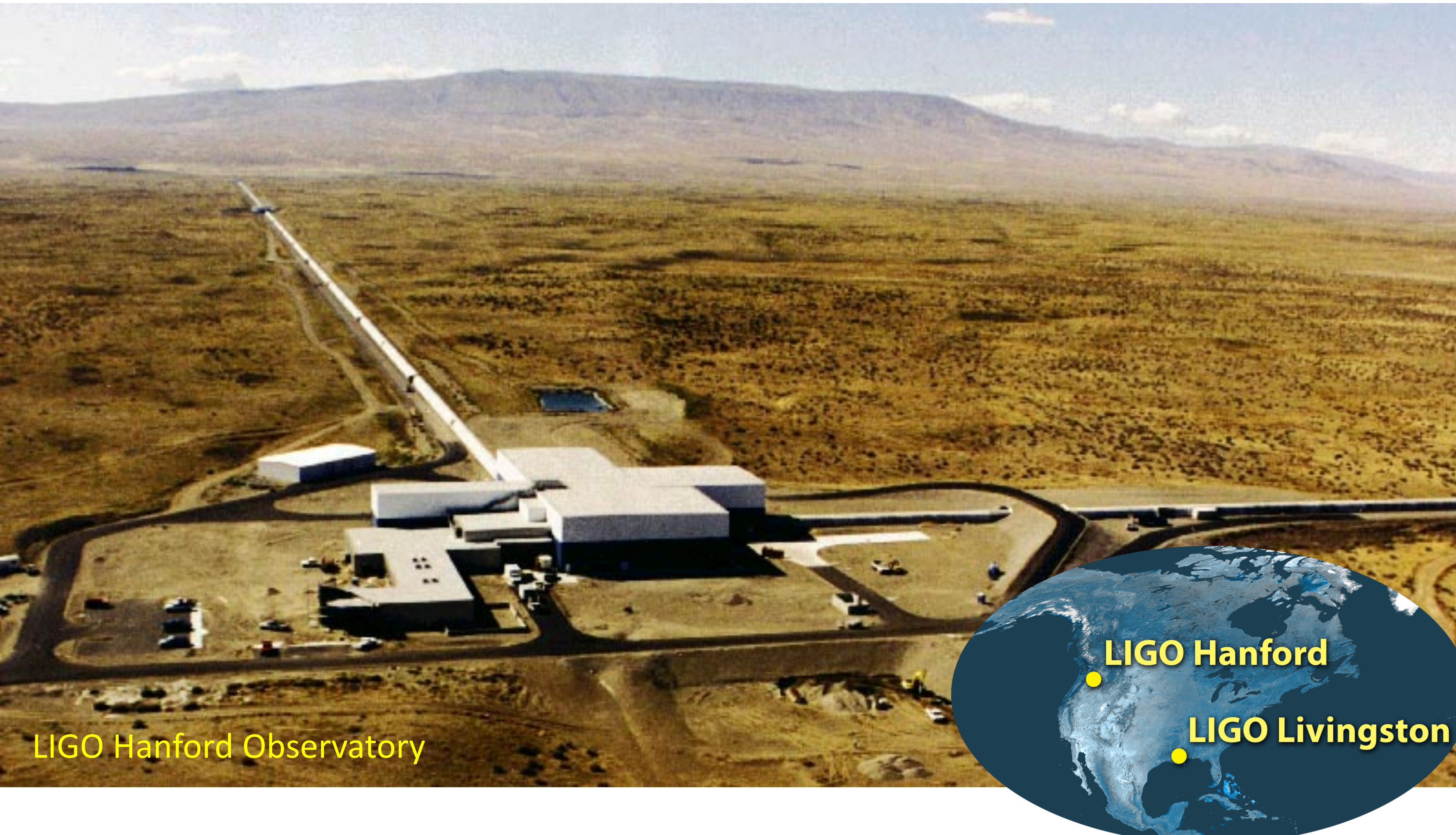
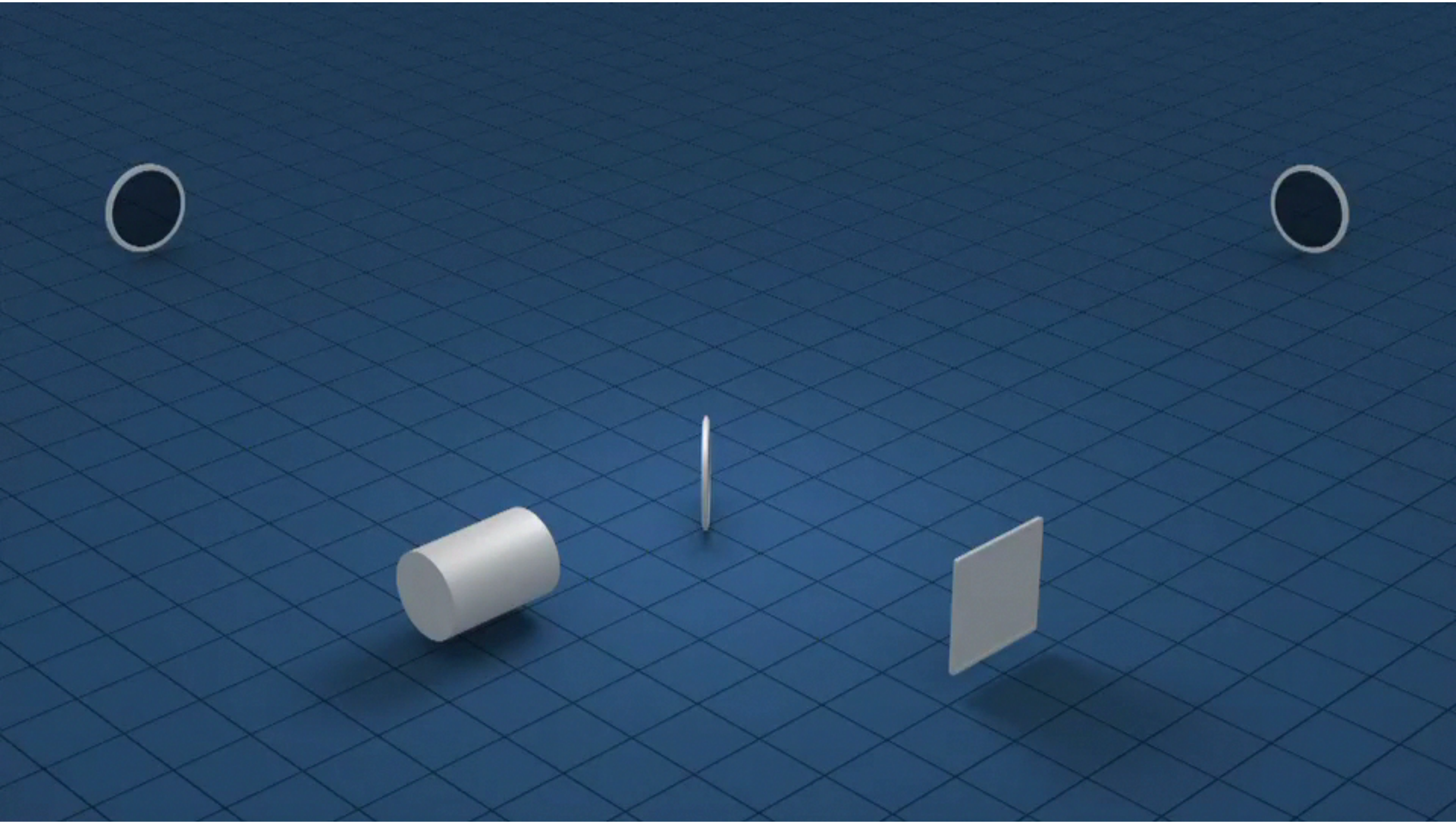

LIGO Hanford Observatory
LIGO Hanford
LIGO Livingston



LIGO The Scale of the Challenge

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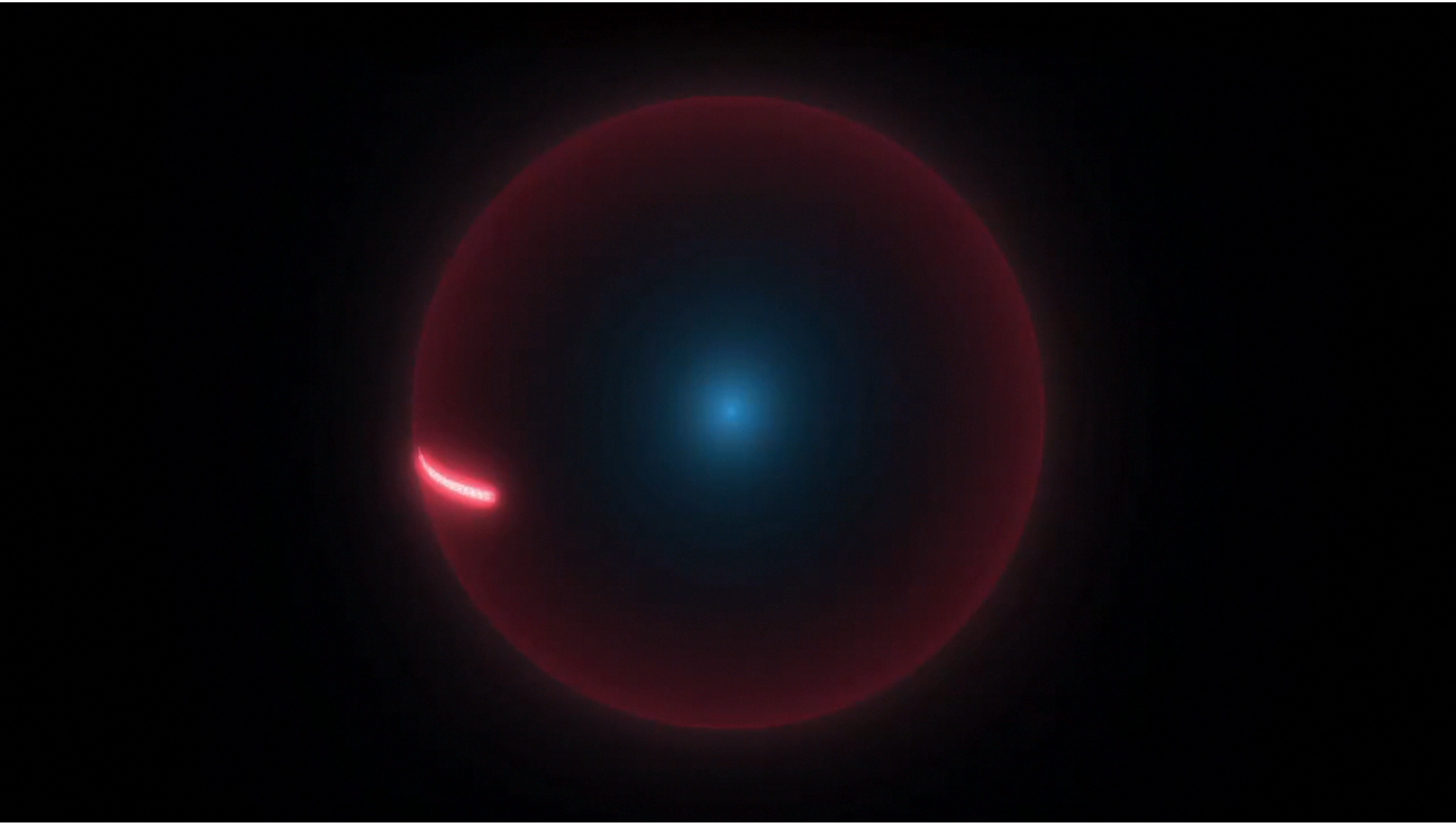
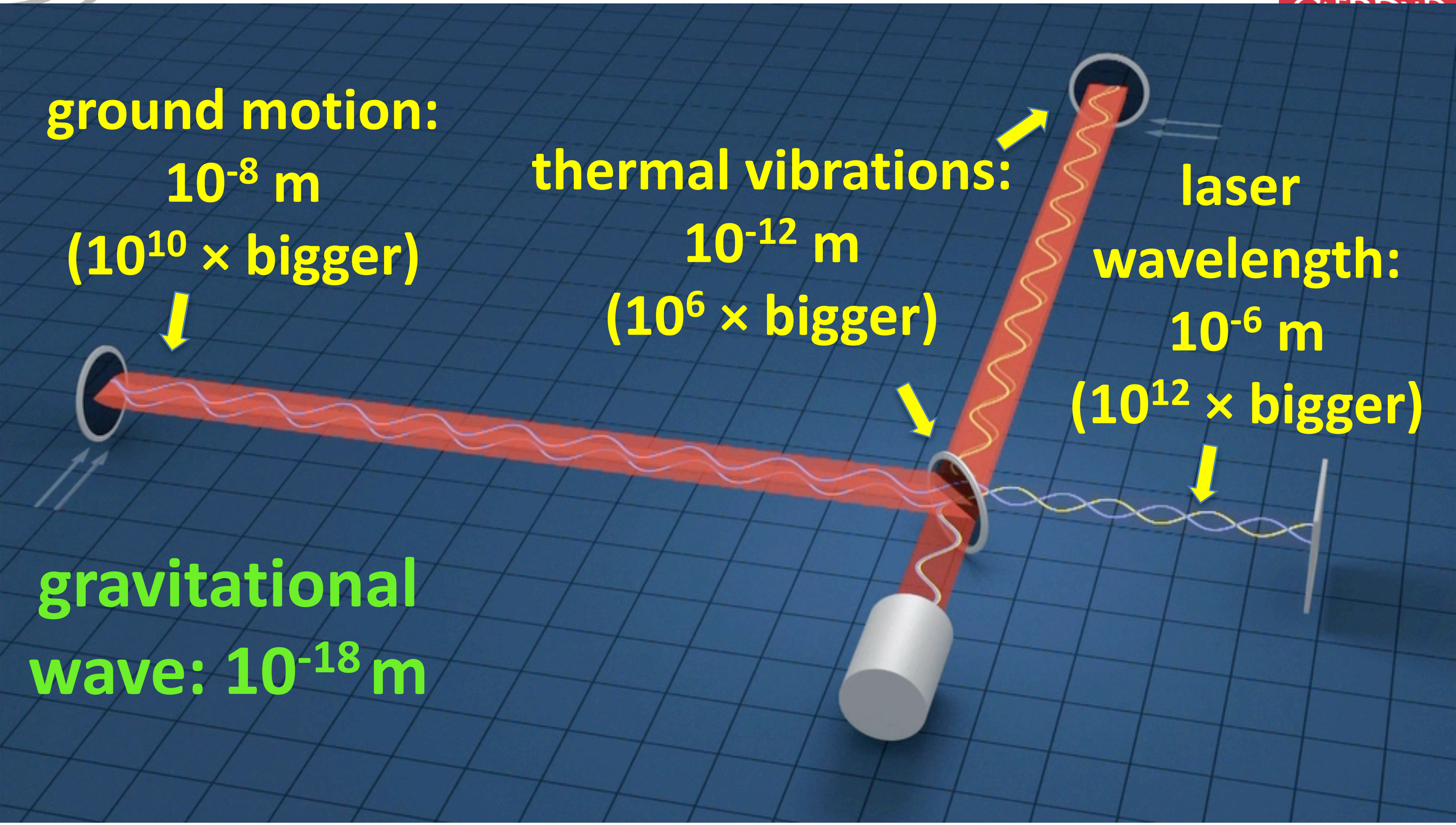
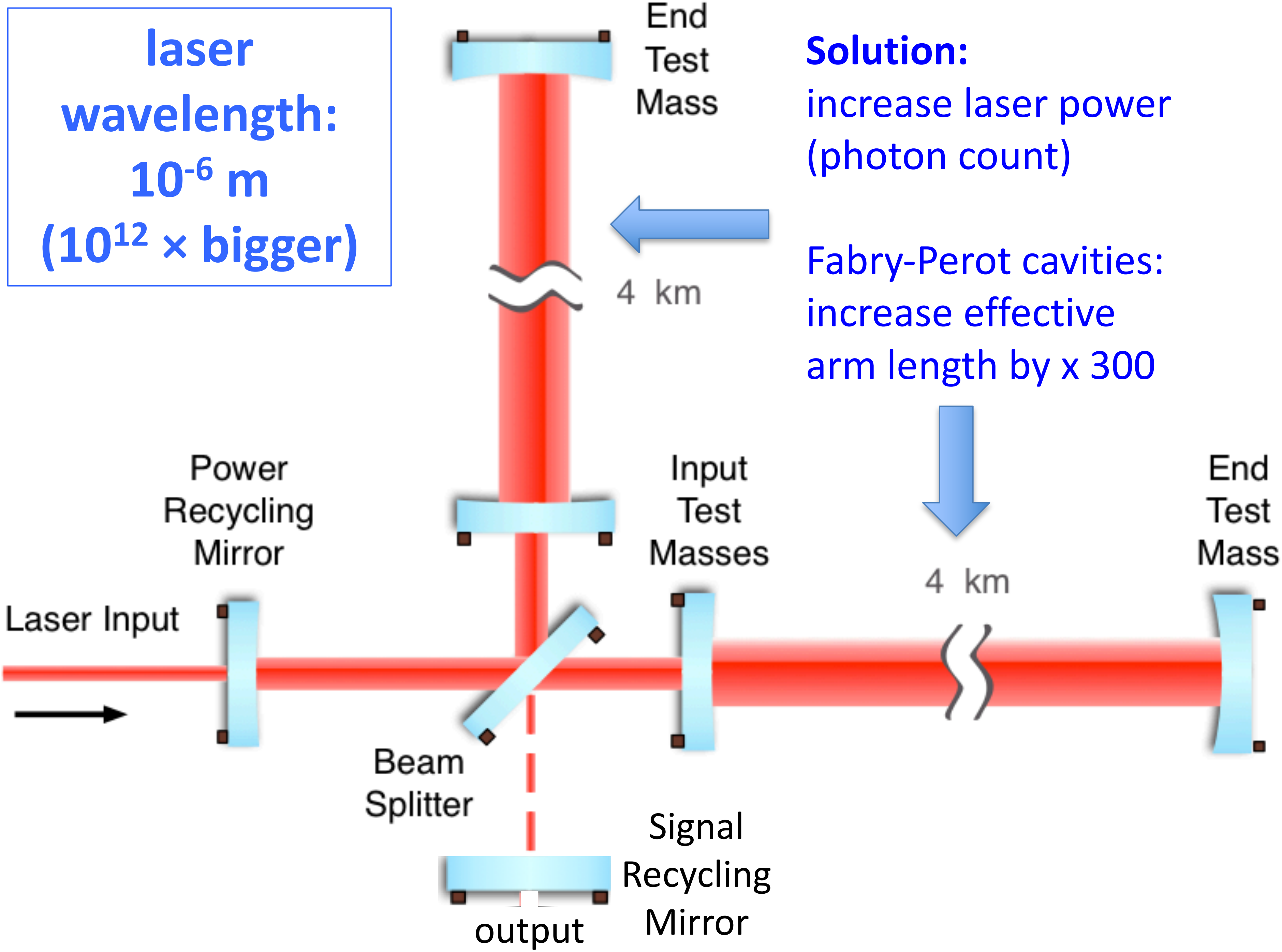
10^{-22} change in length of a LIGO arm: 10^{-18} m

LIGO The Scale of the Challenge




ground motion: 10^{-8} m (10^{10} × bigger)
thermal vibrations: 10^{-12} m (10^{6} × bigger)
laser wavelength: 10^{-6} m (10^{12} × bigger)
gravitational wave: 10^{-18} m

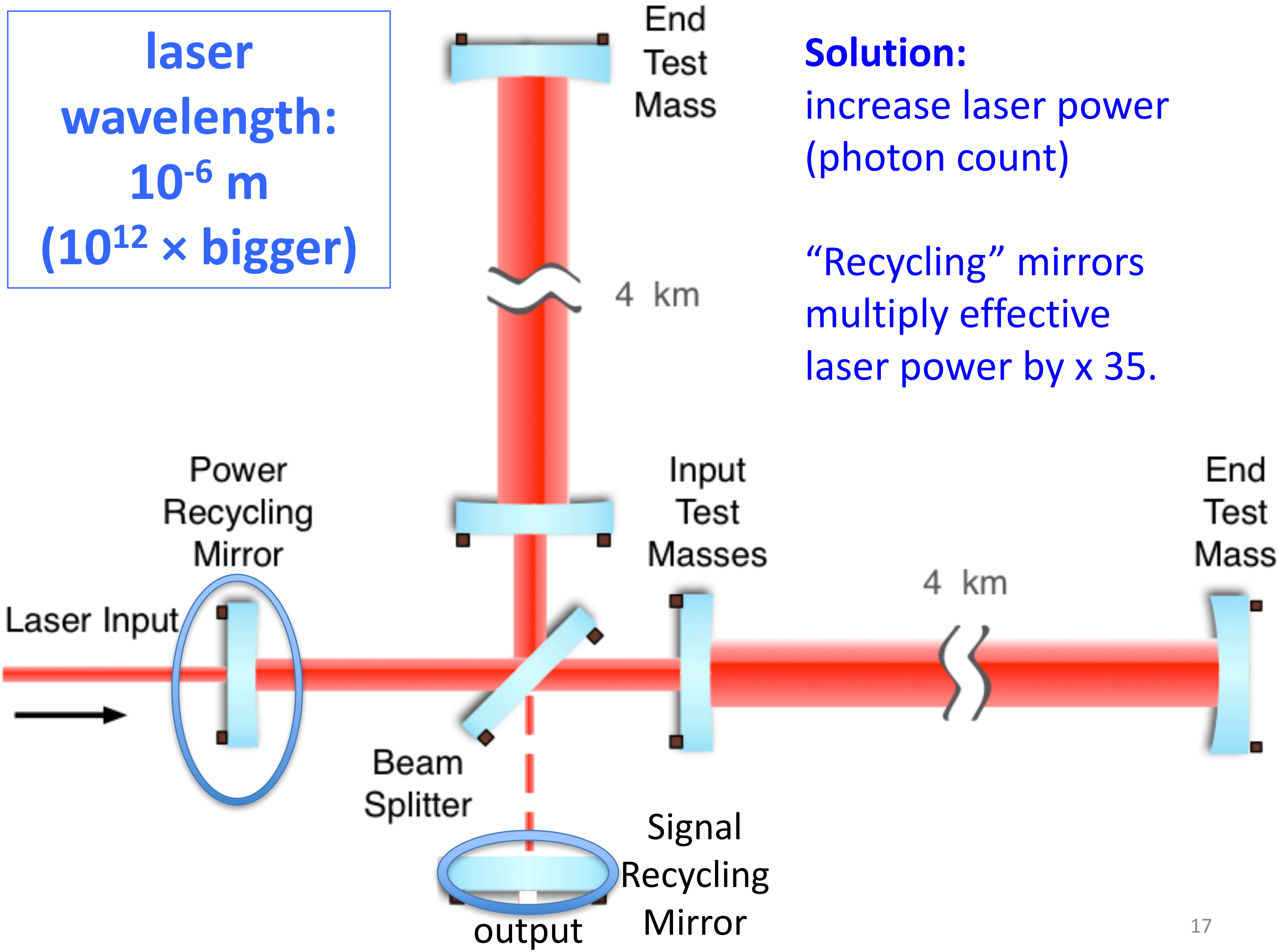

laser wavelength: 10^-6 m
(10^12 × bigger)
Solution:
increase laser power (photon count)
Fabry-Perot cavities: increase effective arm length by x 300
End Test Mass
4 km
Power Recycling Mirror
Input Test Masses
End Test Mass
4 km
Laser Input
Beam Splitter
Signal Recycling Mirror
output

laser wavelength: 10^{-6} m (10^{12} × bigger)

Solution:

increase laser power (photon count)

"Recycling" mirrors multiply effective laser power by x 35.



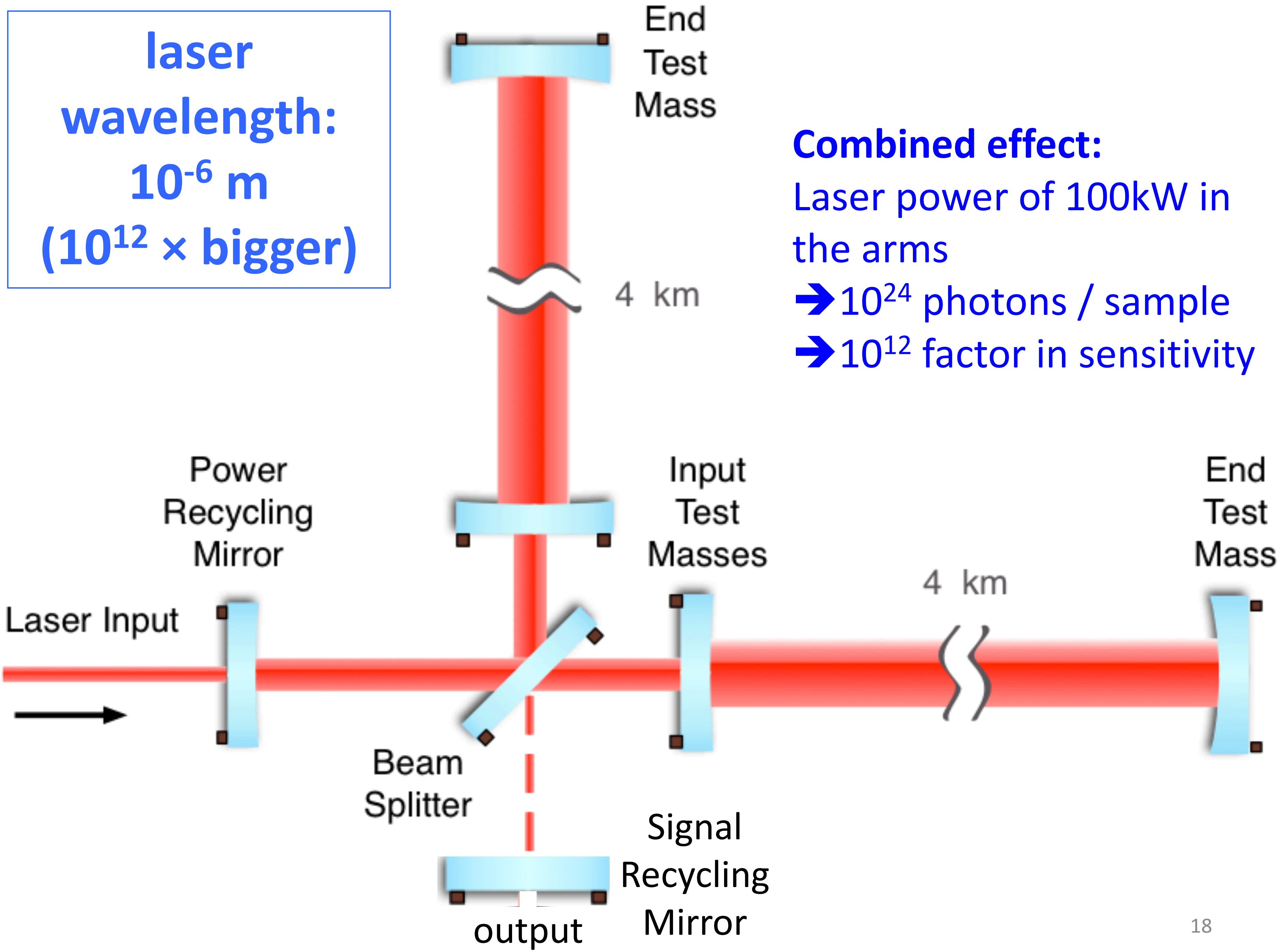
laser wavelength: 10^{-6} m ($10^{12} \times$ bigger)

Combined effect:

Laser power of 100kW in the arms

→ 10^{24} photons / sample

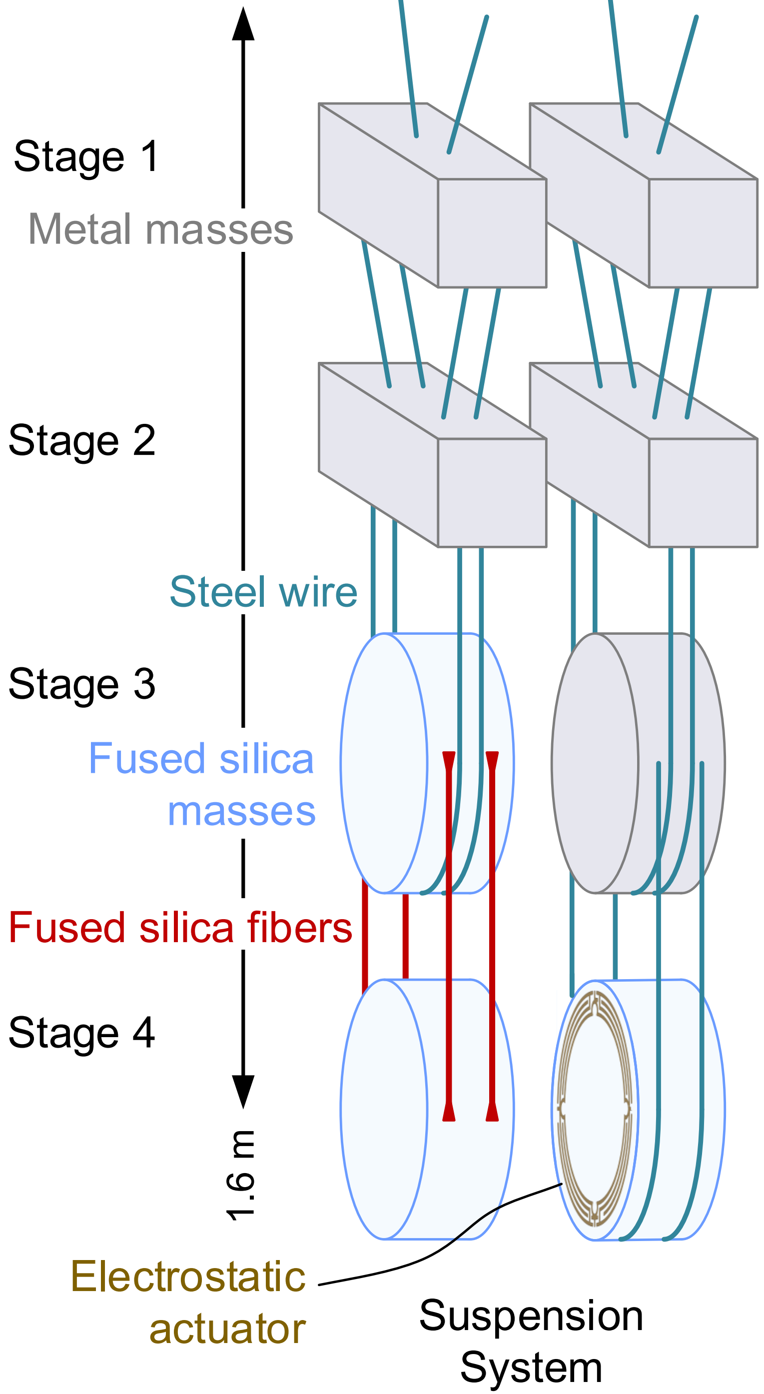
→ 10^{12} factor in sensitivity

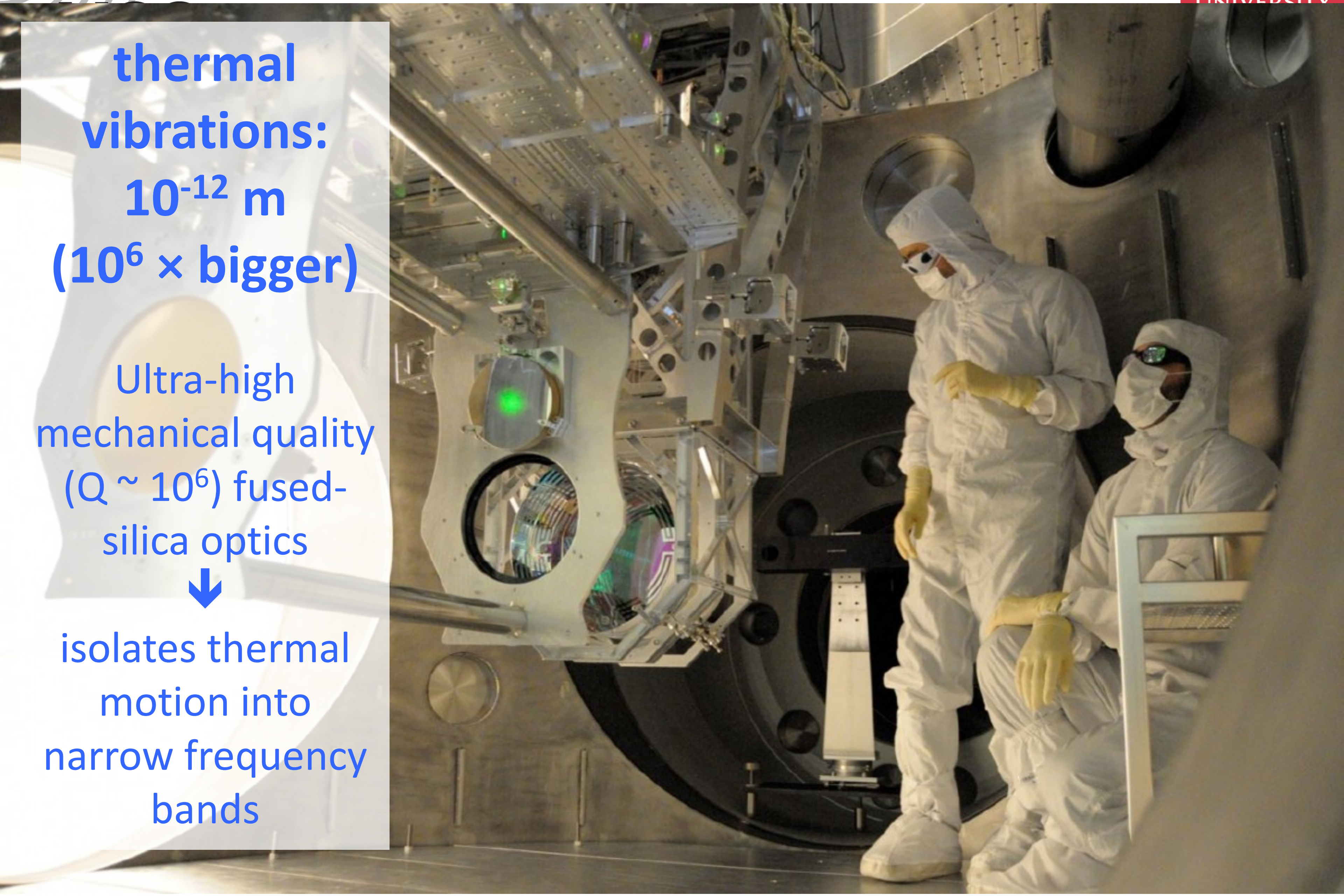


LIGO

ground motion:
10^{-8} m
(10^{10} × bigger)

Quadruple pendulum suspension system: **10^7**
+
Active seismic isolation: **10^3**





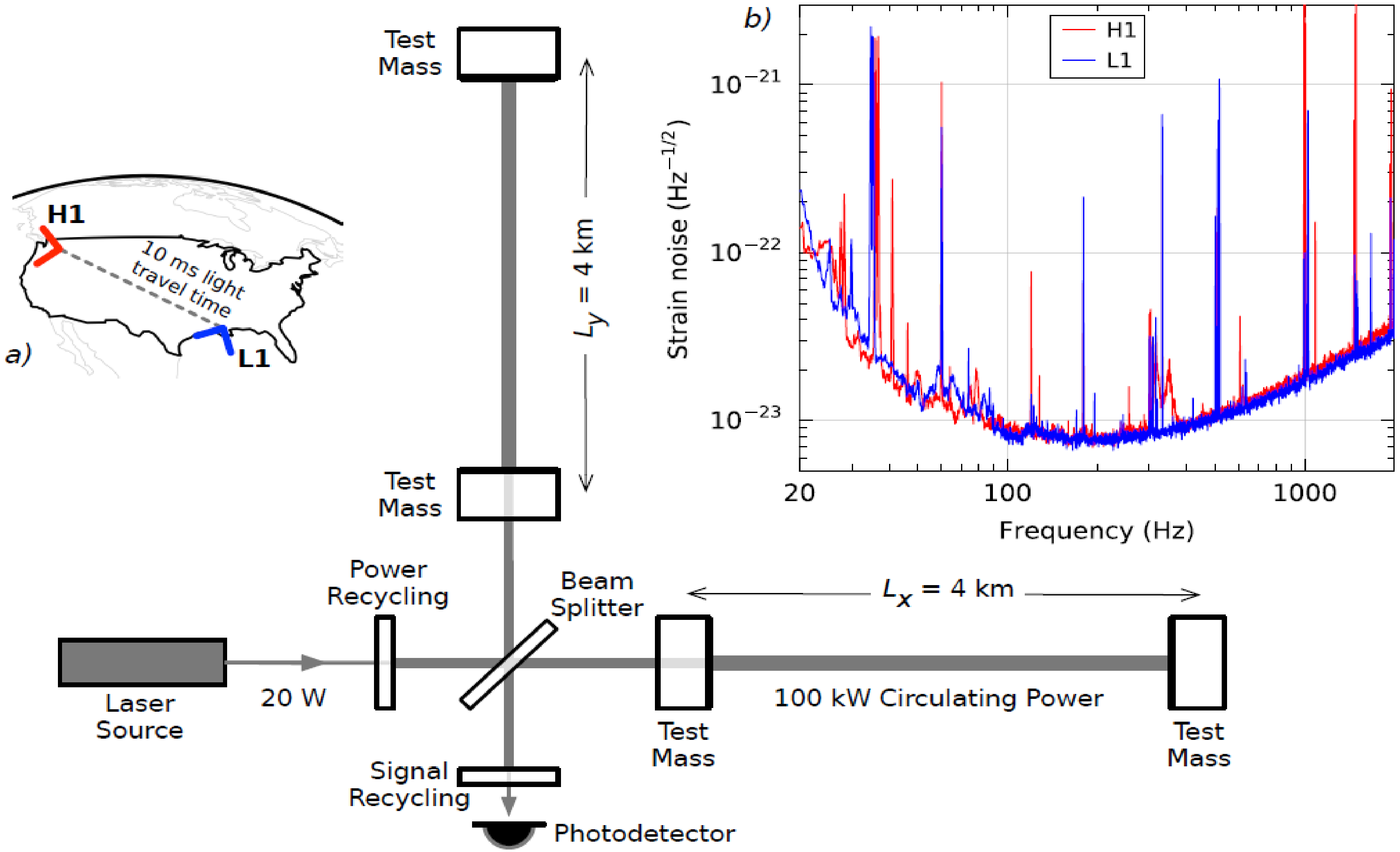
thermal vibrations: 10^{-12} m ($10^6 \times$ bigger)

Ultra-high mechanical quality ($Q \sim 10^6$) fused-silica optics

↓

isolates thermal motion into narrow frequency bands

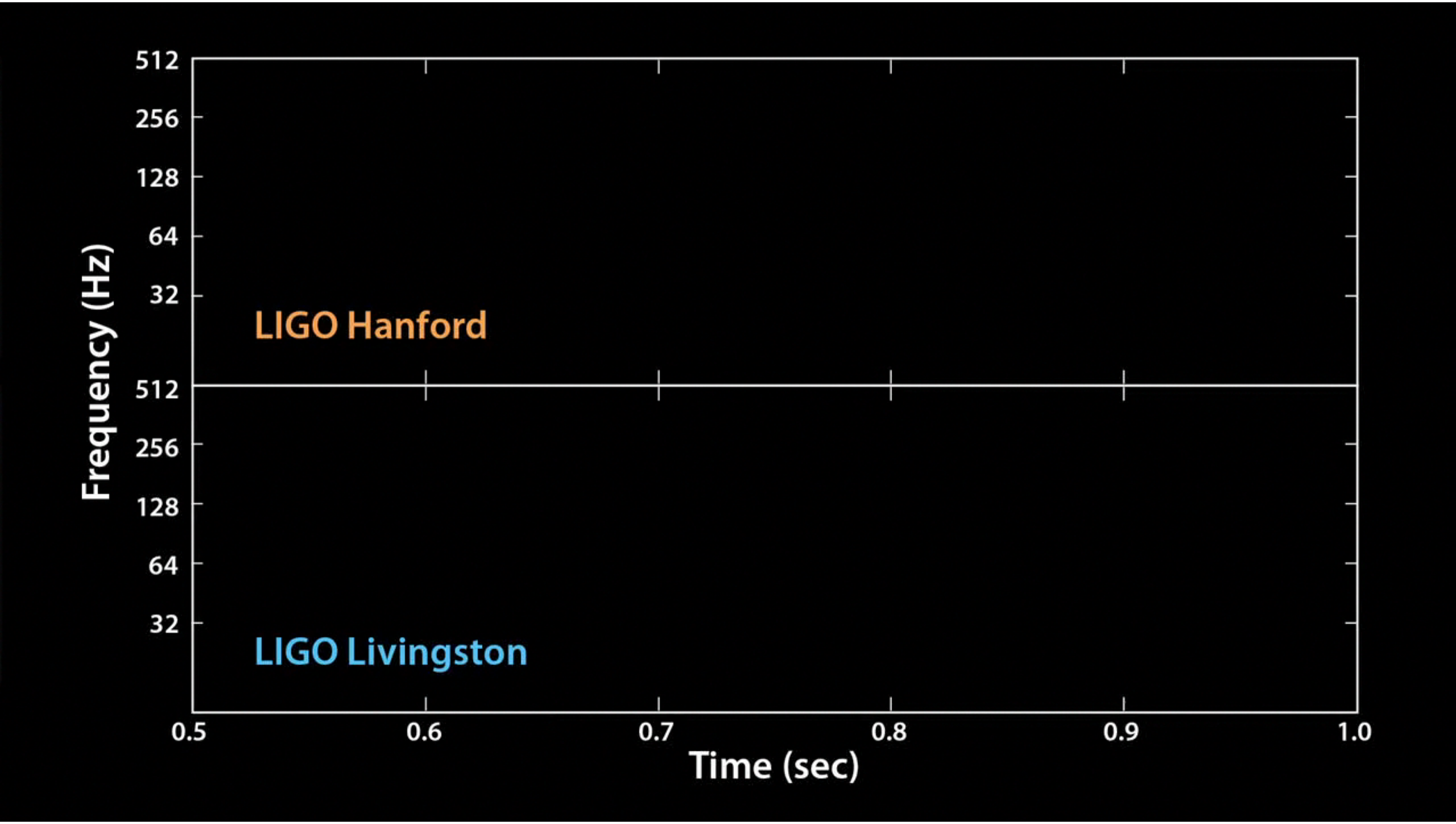
O1 Data Taking



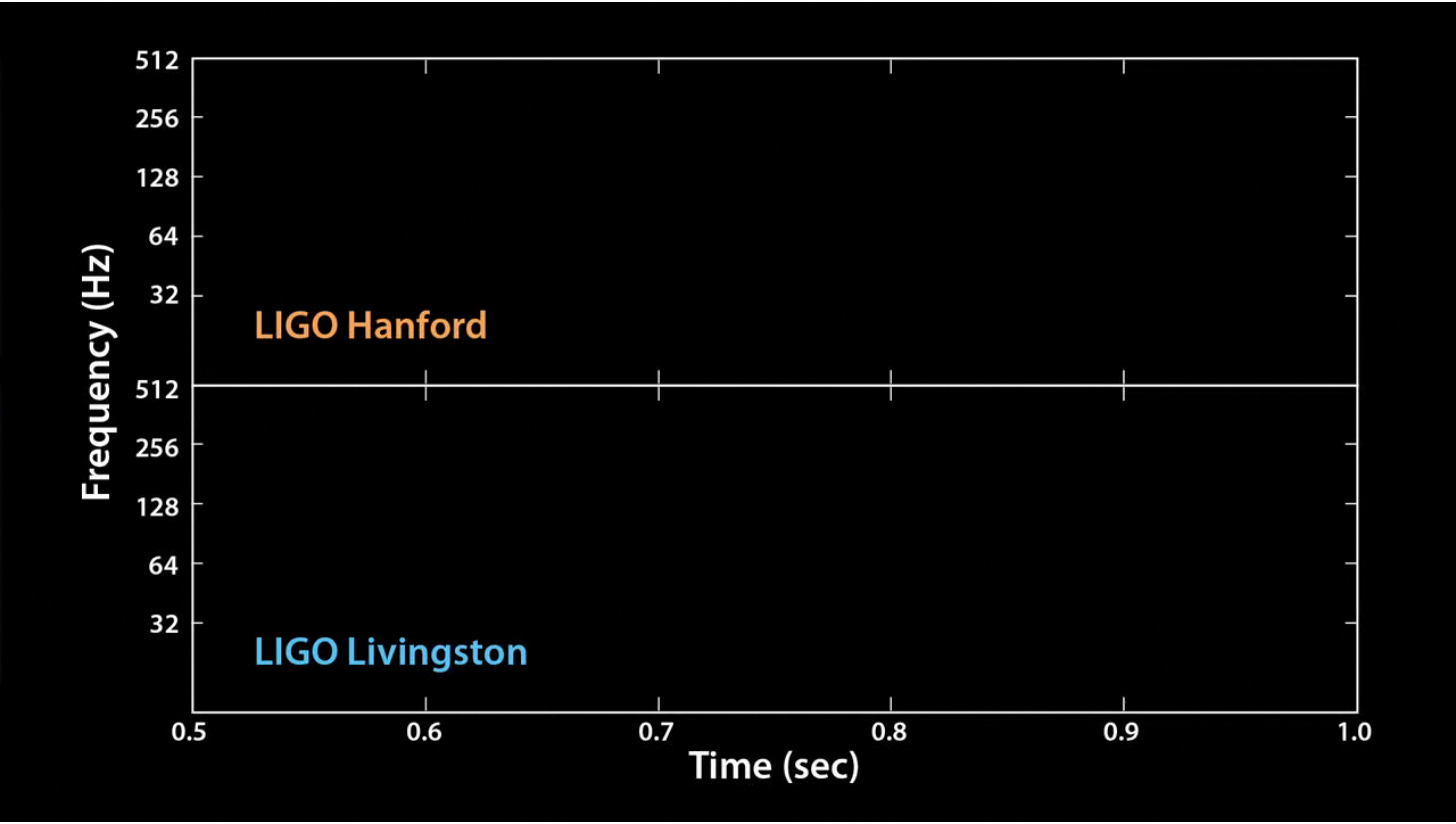
O1 Data Taking

- In September 2015, we were in the final stages of preparation for first Advanced LIGO data run (O1).
- The very last step is a short “Engineering Run,” during which on Sept 14 our online monitor recorded GW150914.
- We identified the signal within 3 minutes
- We responded by starting the data run officially, keeping all settings fixed and ran for 16 live days coincidence time (long enough to assess background levels, etc)
- First GW announcement reported on that data.
- O1 data taking continued until 12 Jan 2016

GW150914

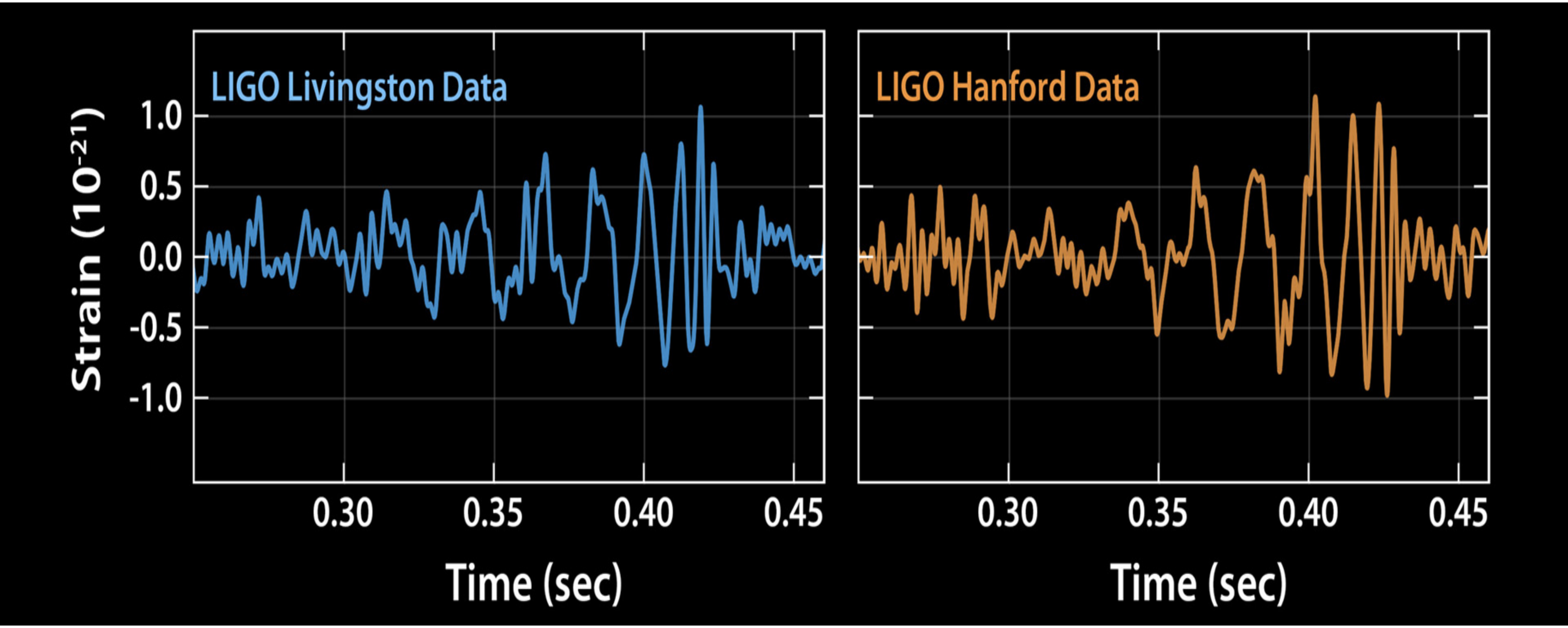

512
256
128
64
32
LIGO Hanford
512
256
128
64
32
LIGO Livingston
Frequency (Hz)
0.5
0.6
0.7
0.8
0.9
1.0
Time (sec)

GW150914



512
256
128
64
32
LIGO Hanford
512
256
128
64
32
LIGO Livingston
Frequency (Hz)
0.5
0.6
0.7
0.8
0.9
1.0
Time (sec)

GW 150914



LIGO Livingston Data
LIGO Hanford Data
Strain (10⁻²¹)
1.0
0.5
0.0
-0.5
-1.0
0.30
0.35
0.40
0.45
Time (sec)
Time (sec)

LIGO

GW 150914

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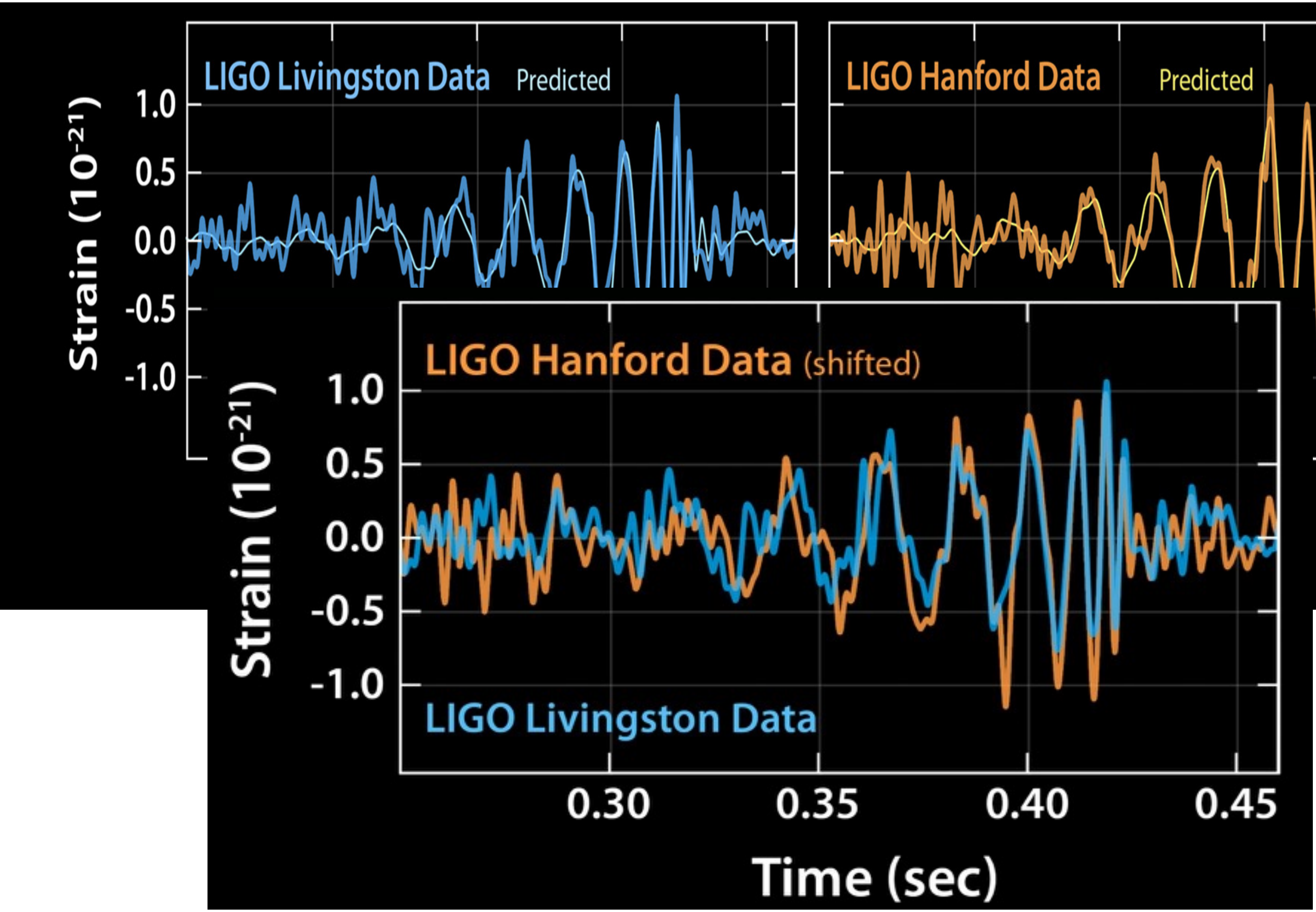
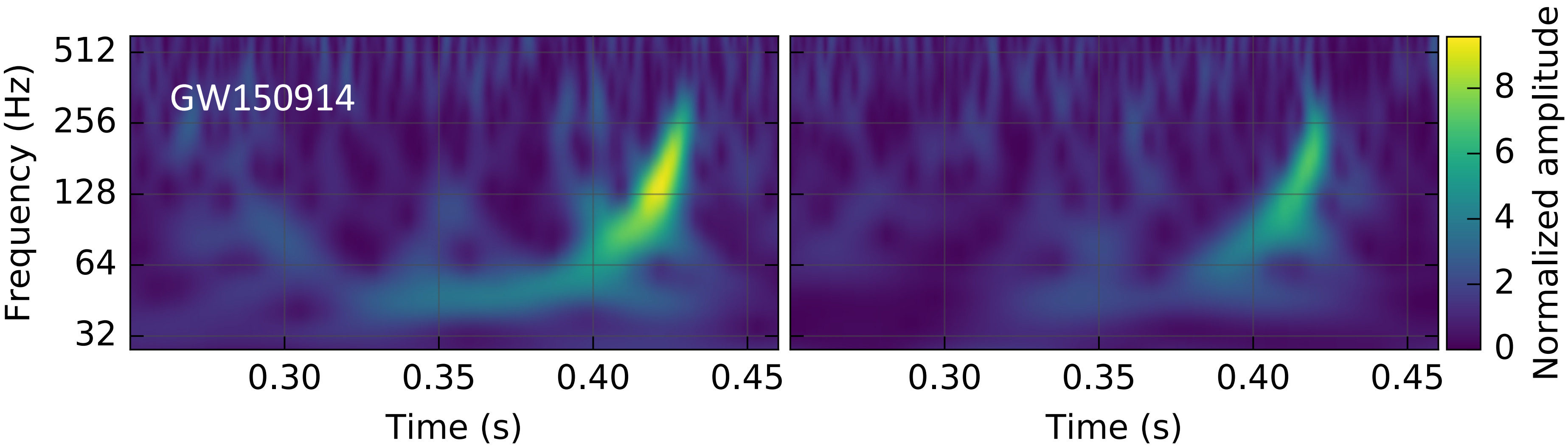
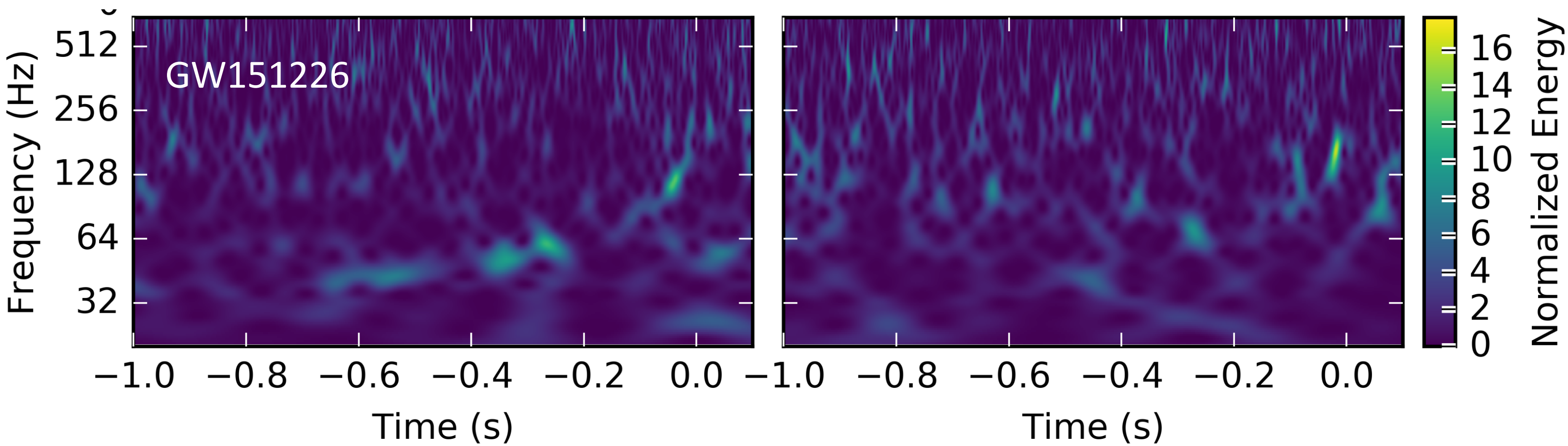


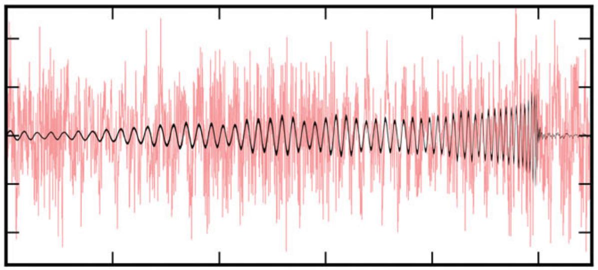
Identifying the signals





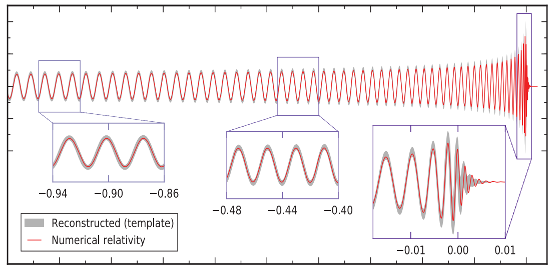
Matched filtering

Data



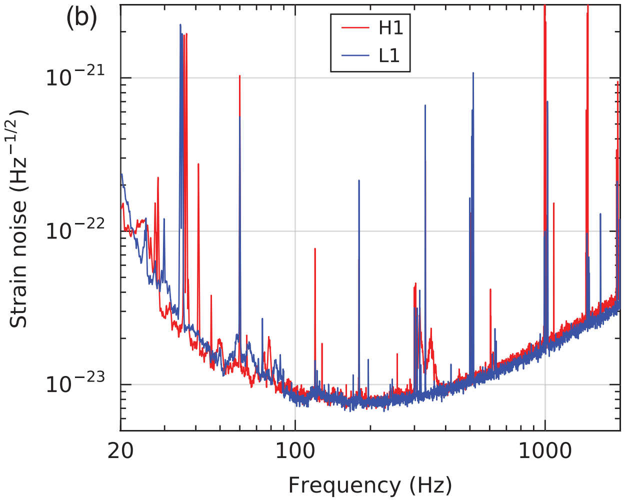
X

Waveform



/

Sensitivity



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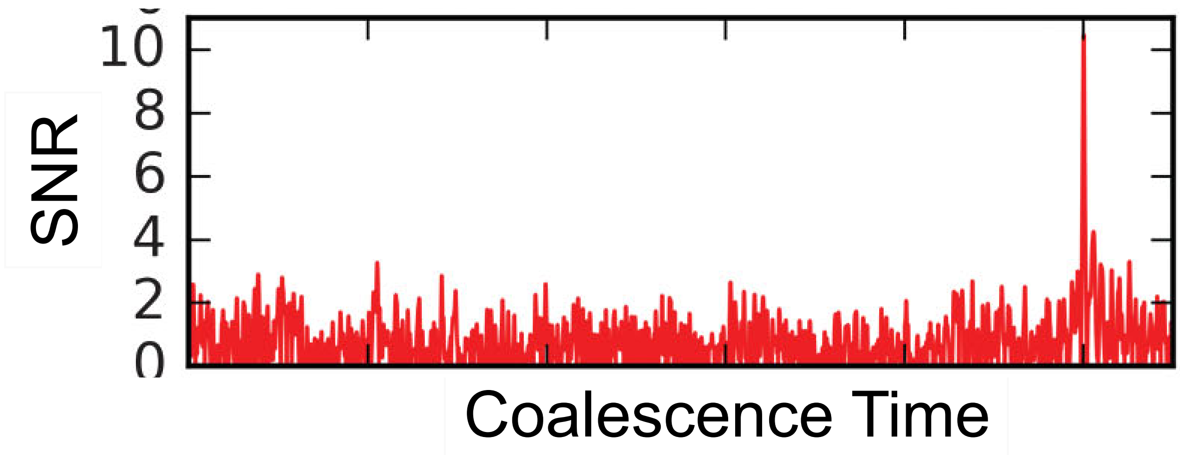
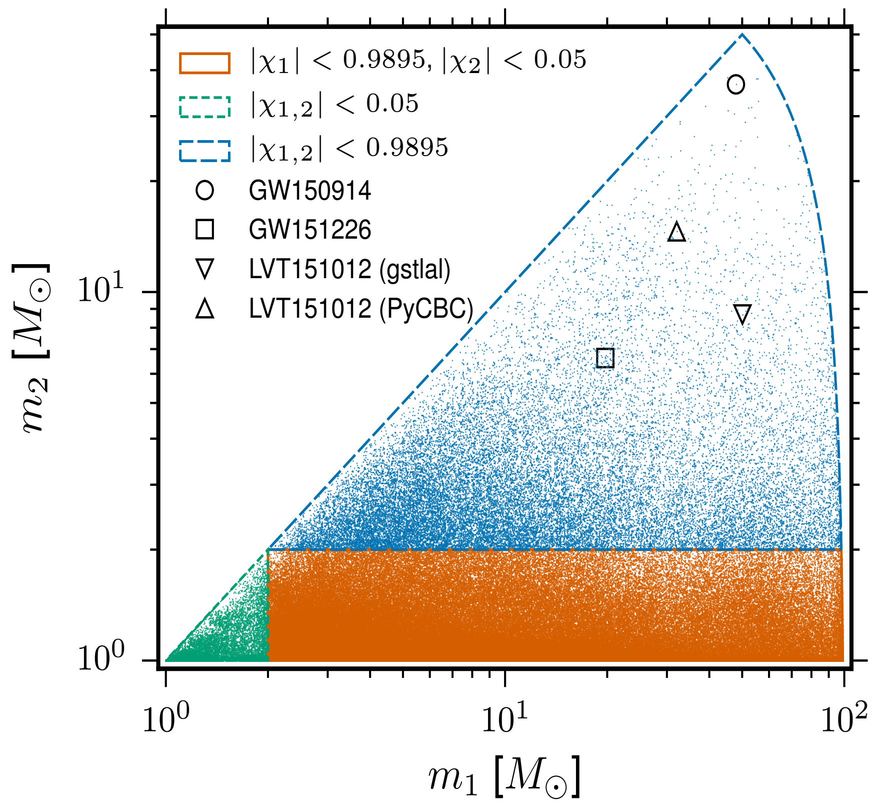
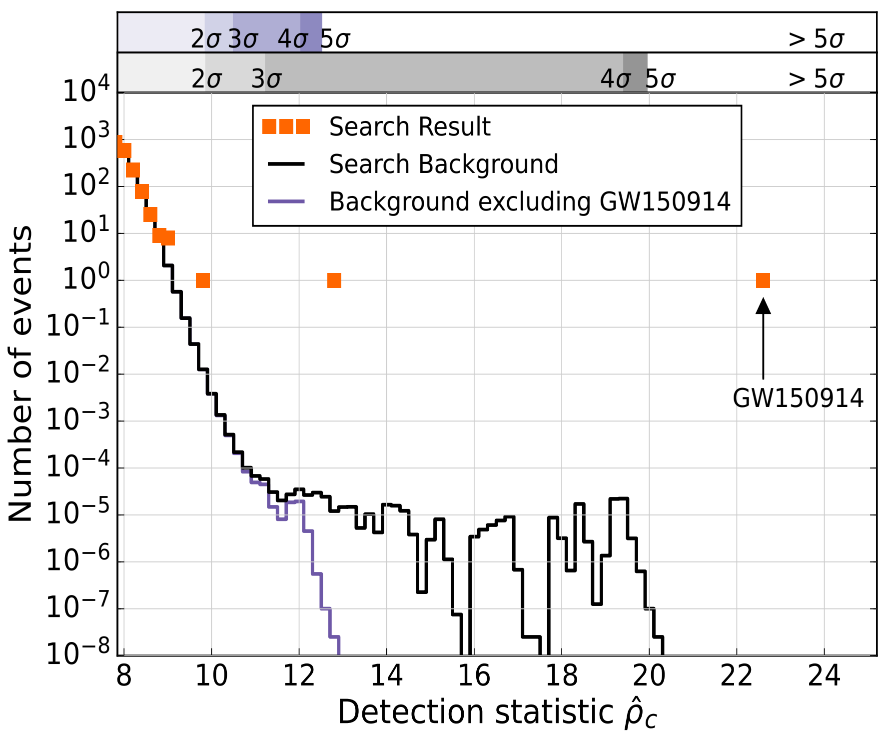


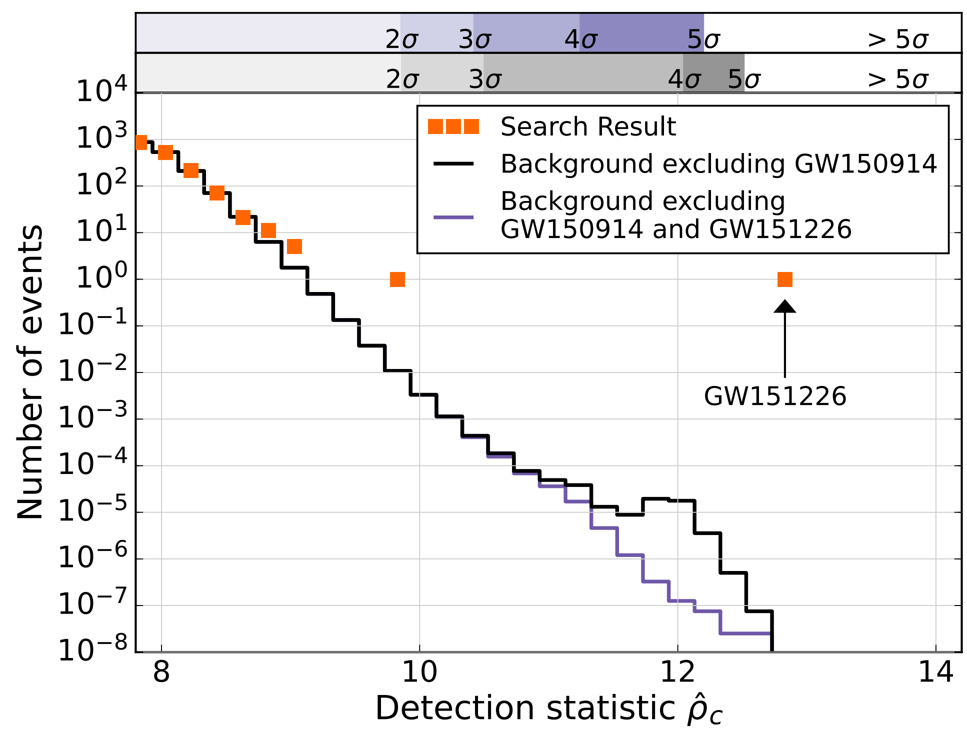
Binary Merger Search

- Use known waveforms to search for binary signals.
- Calculate Signal to Noise Ratio (SNR), ρ(t), identify maxima, and reweight by a χ^2 consistency measure.
- Require coincidence between detectors within 15 msec.
- Detection statistic: quadrature sum of the signal to noise in each detector.
- Background: Time shift by multiples of 0.1 seconds and repeat search.



Statistical Significance

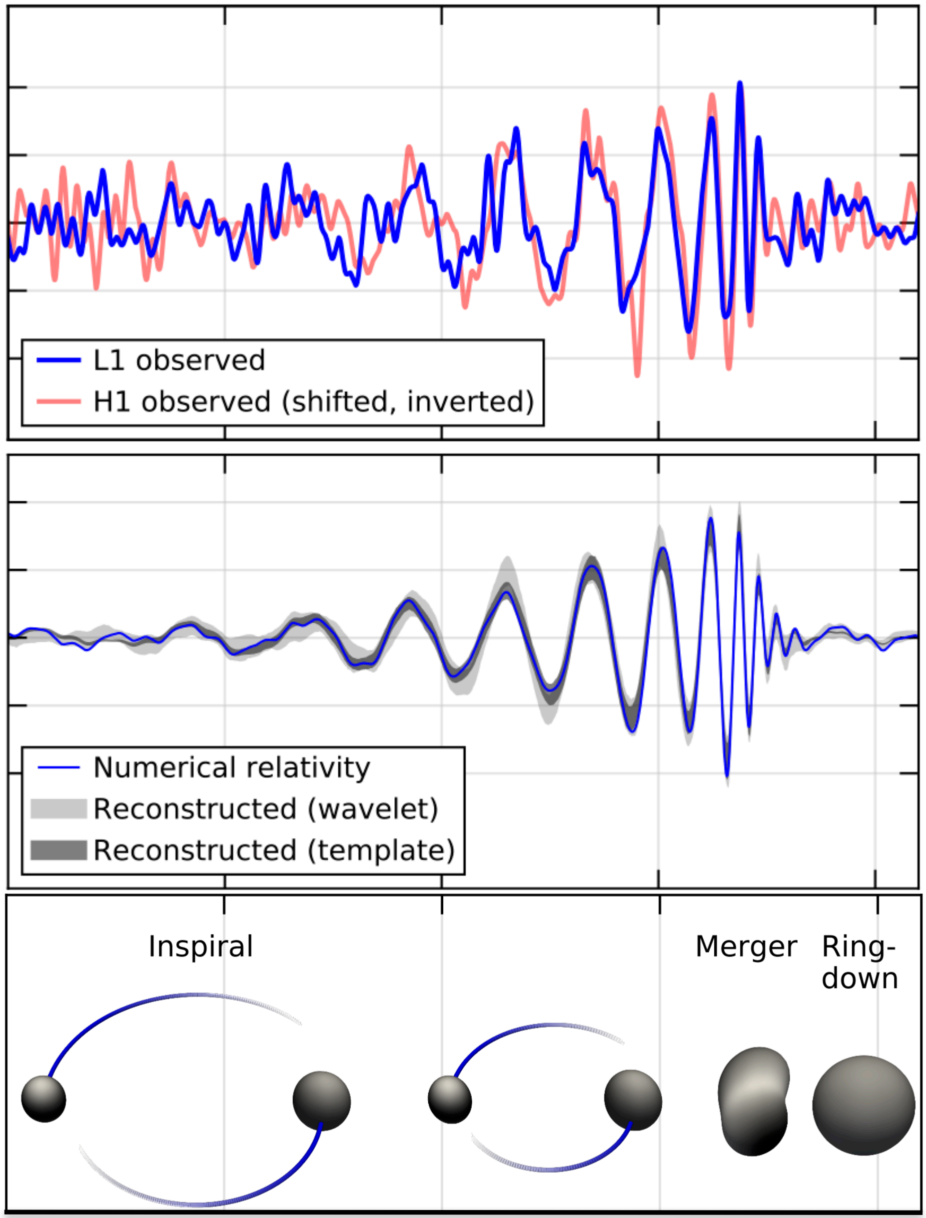




A black hole binary

- Orbits decay due to emission of gravitational waves

- **Leading order** determined by "chirp mass"

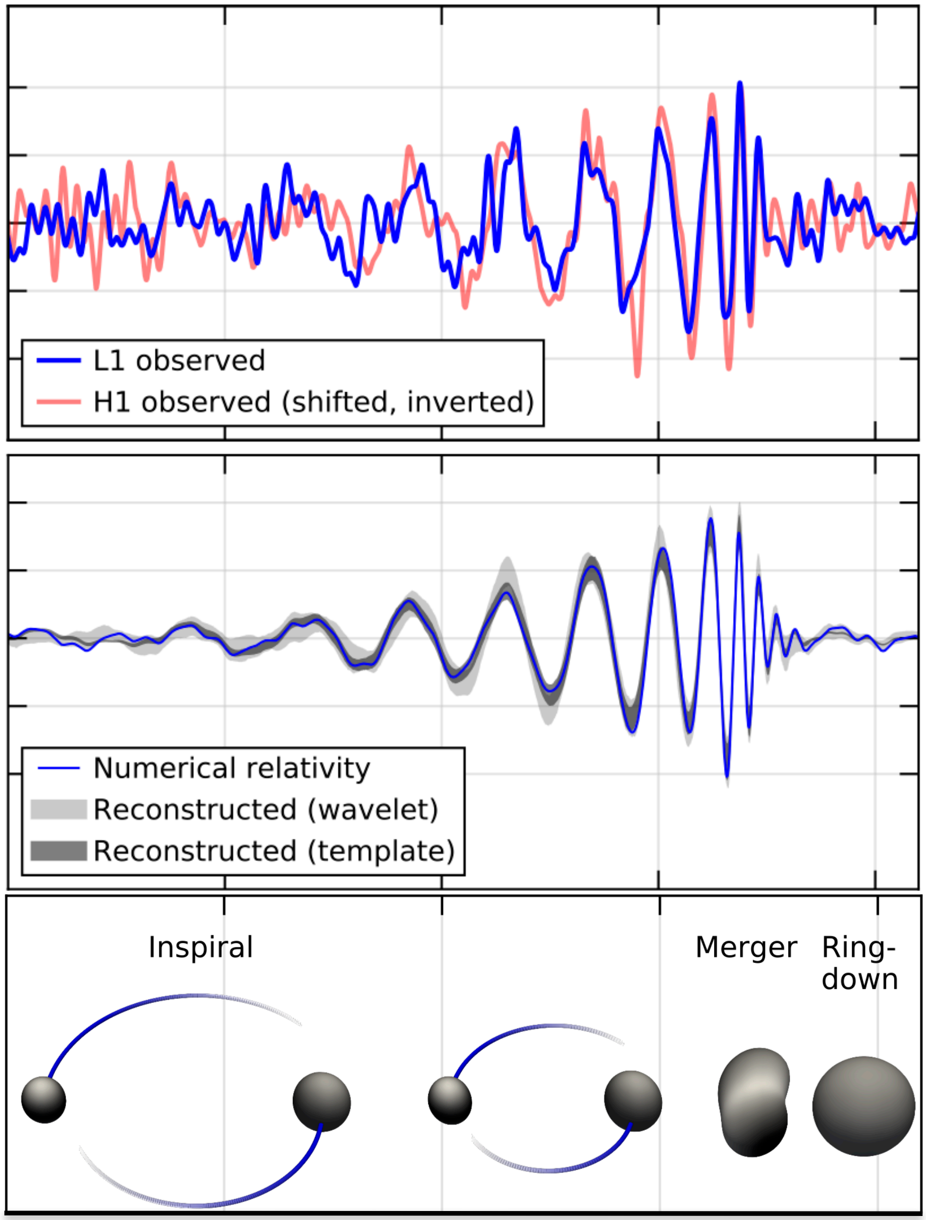


$$\mathcal{M} = \frac{(m_1 m_2)^{3/5}}{M^{1/5}} \simeq \frac{c^3}{G}\left[\frac{5}{96}\pi^{-8/3} f^{-11/3} \dot{f}\right]^{3/5}$$

A black hole binary

- Binary is at least sixty times as massive as the sun.

- Bodies are in orbit until centres are separated by a few hundred km.



LIGO

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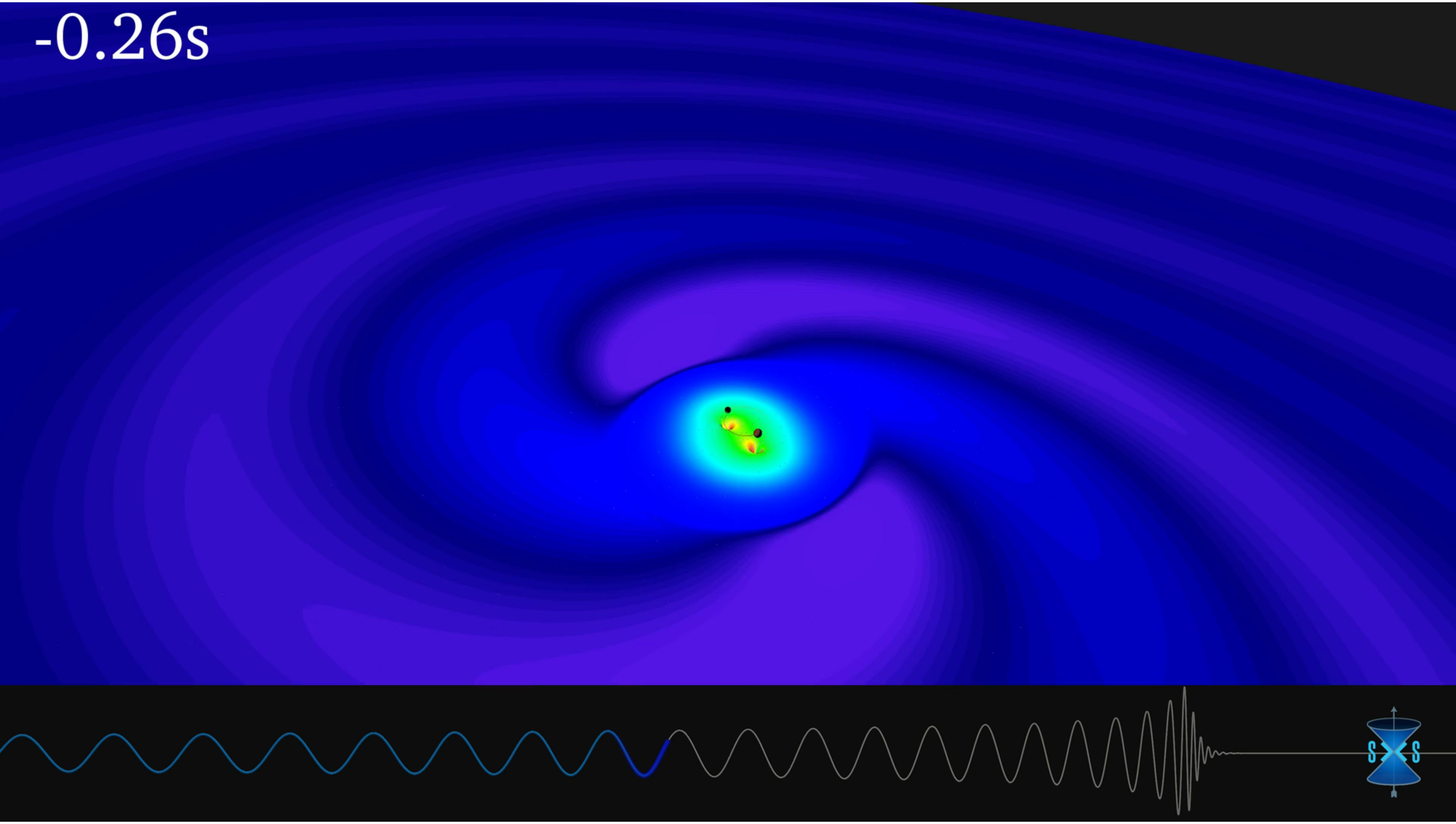
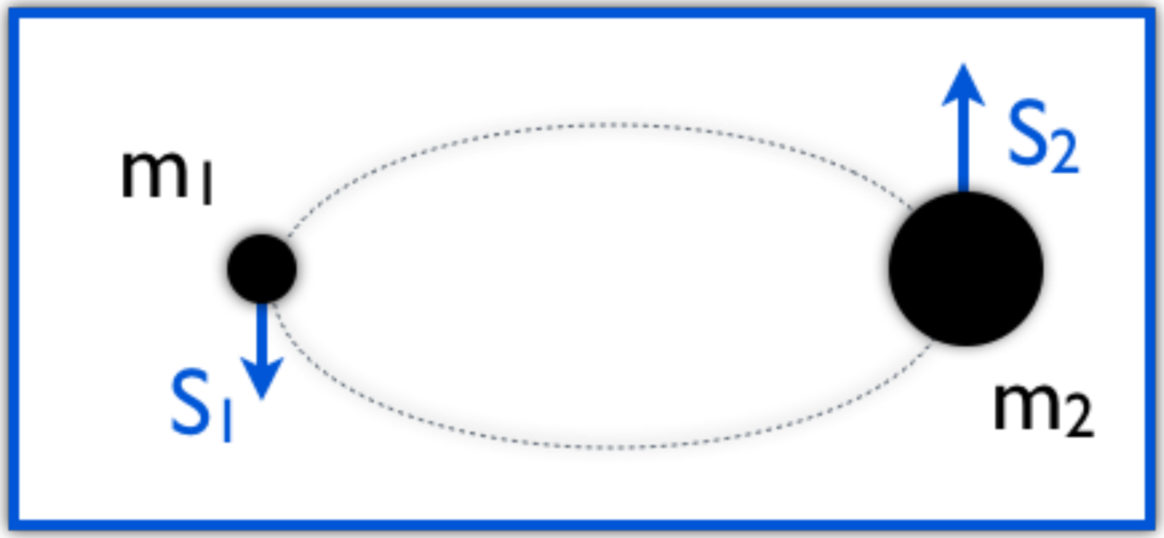


What affects the waveform?

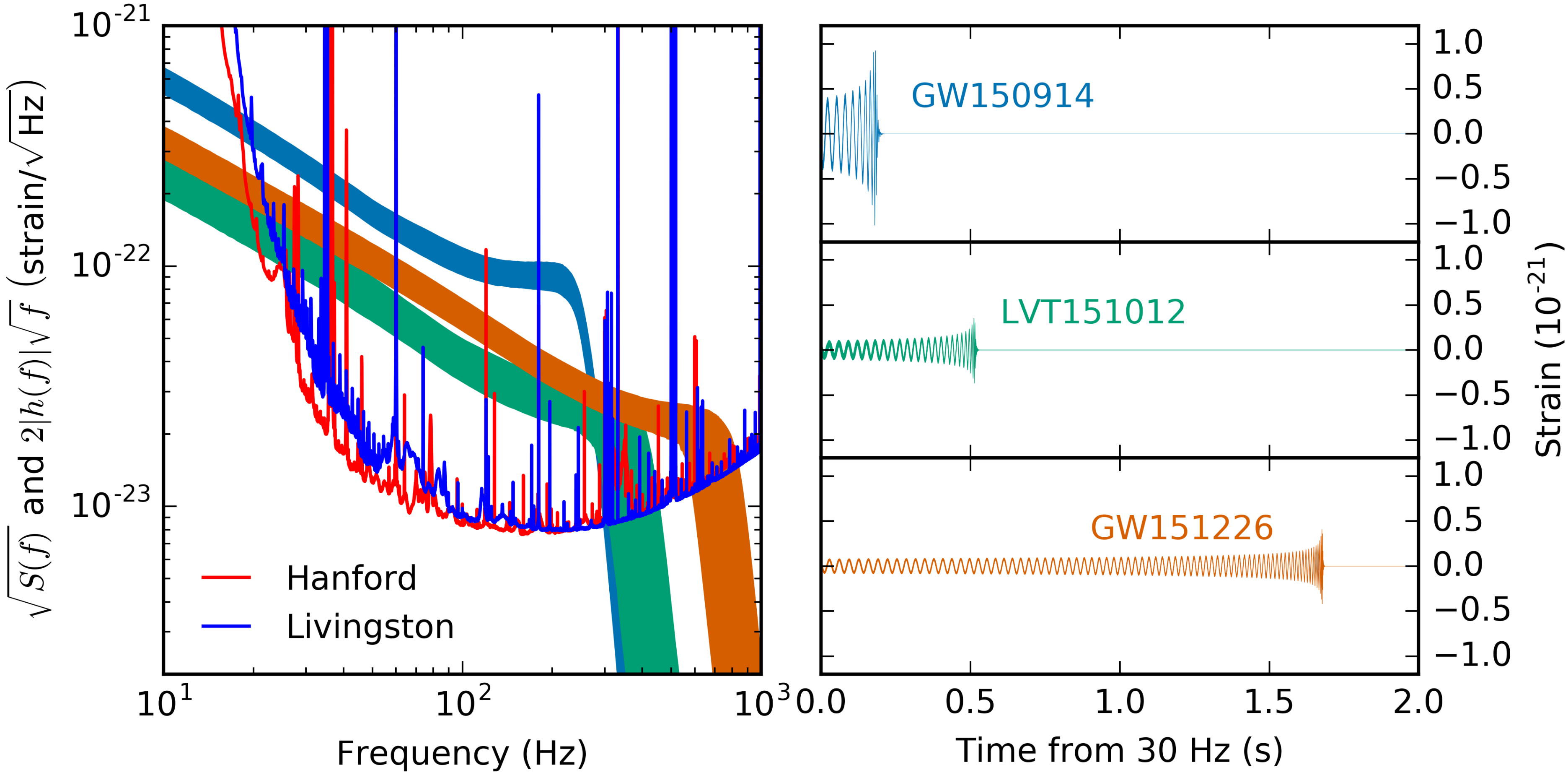
- Total mass -> Change in time scale
- Mass ratio and spins
  -> Change in amplitude / frequency evolution
  ("total" spin has the dominant effect)



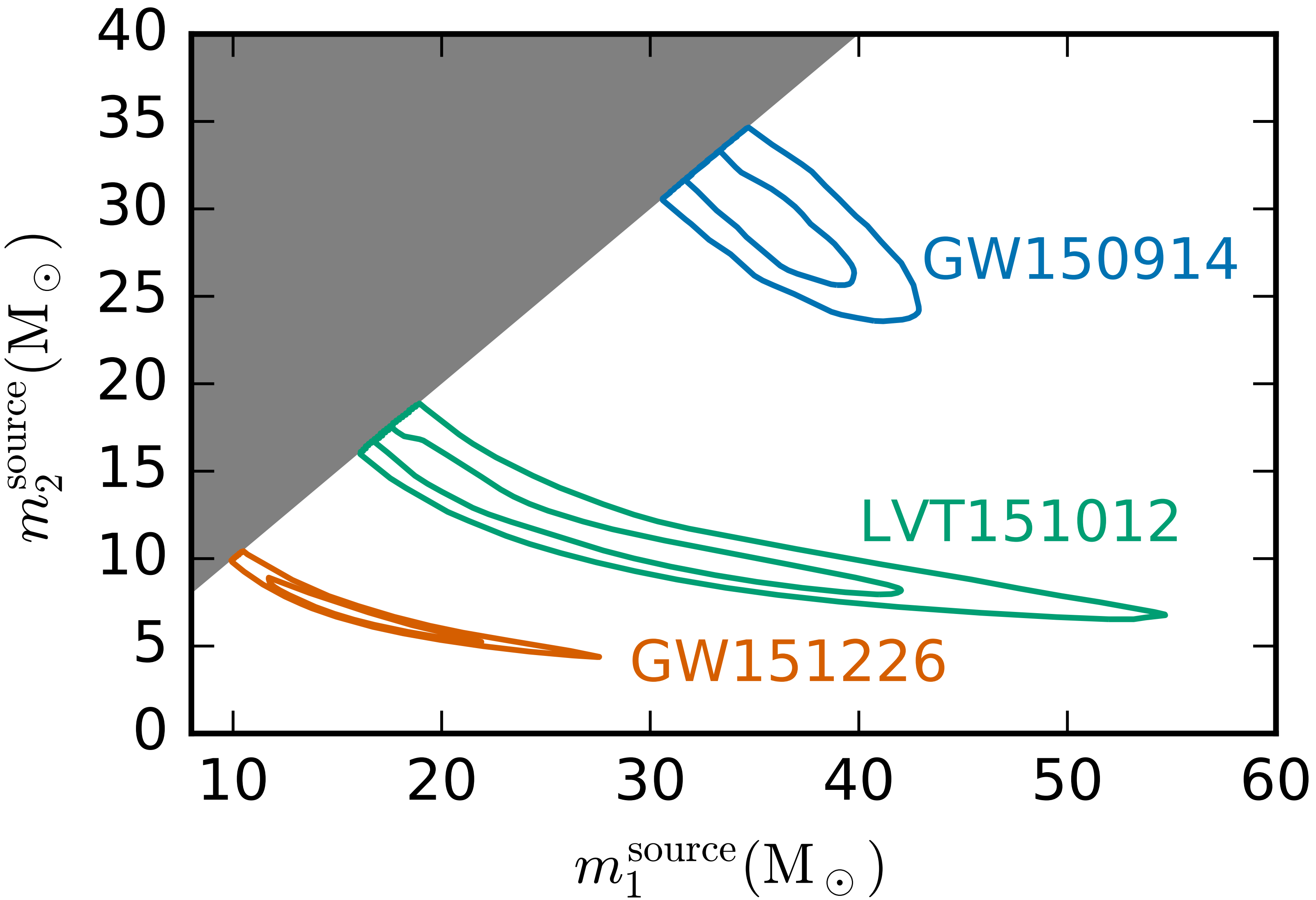
Cosmological effects

- Binary merger signal has a characteristic shape
  - Scales with the mass, M, of the system
- Redshift reduces observed frequencies
  - Indistinguishable from change in mass
    => measure M (1 +z)
- Amplitude scales
  - inversely with the co-moving distance, D_C
  - with the total mass, M
- Directly measure:
  - luminosity distance, $D_L = D_C (1 + z)$
  - Redshifted mass, M (1 + z)

Observed Signals

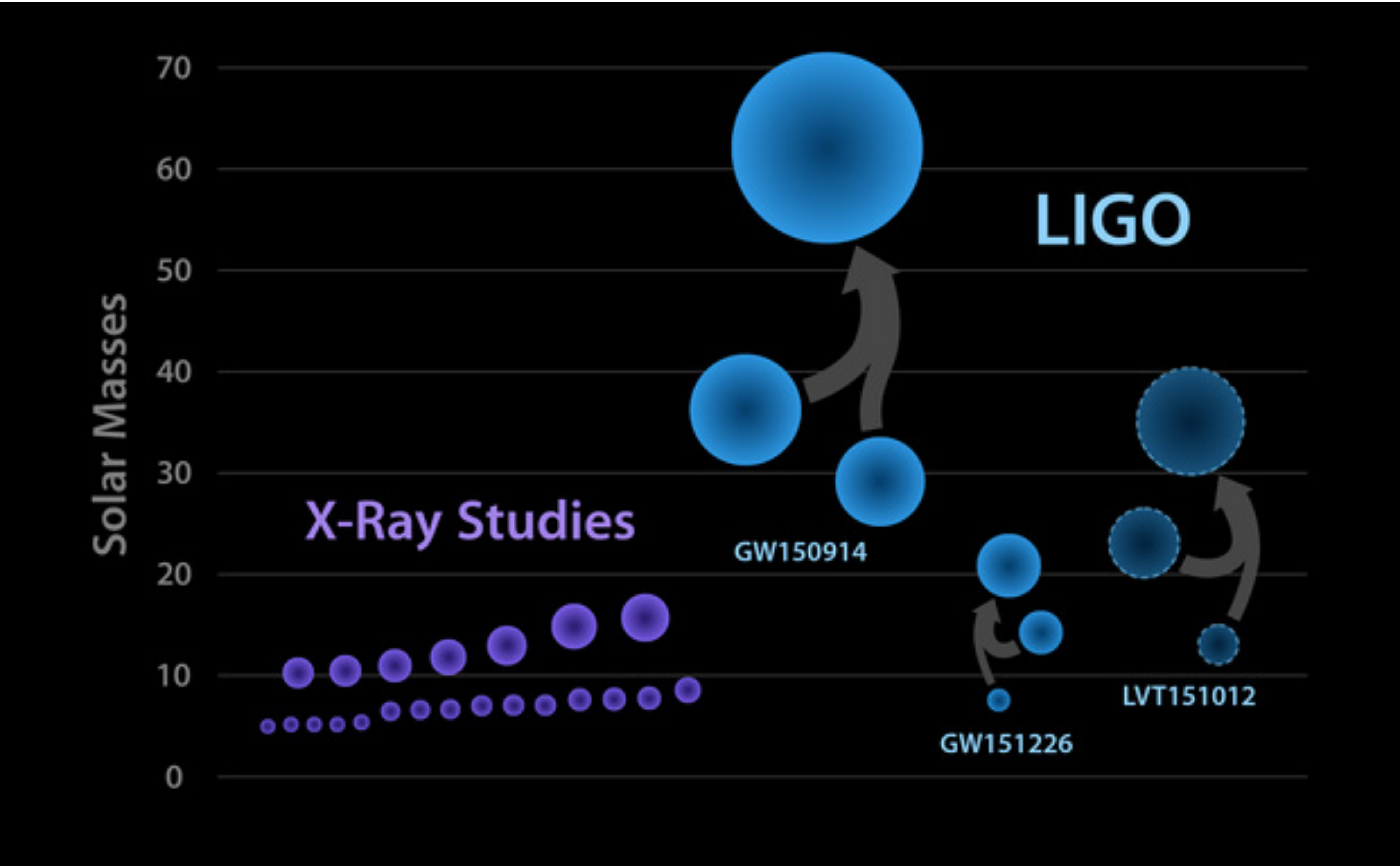


Masses

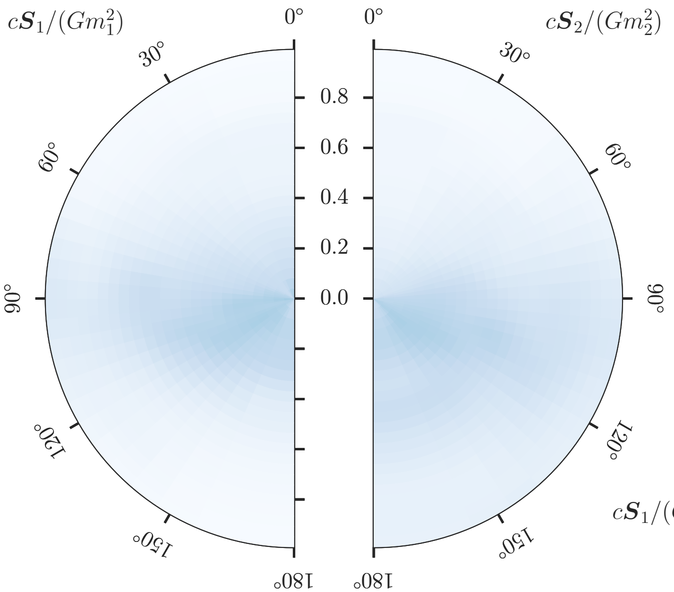


40
35
30
25
20
15
10
5
0
10
20
30
40
50
60
GW150914
LVT151012
GW151226
$m_2^{\mathrm{source}}(\mathrm{M}_\odot)$
$m_1^{\mathrm{source}}(\mathrm{M}_\odot)$

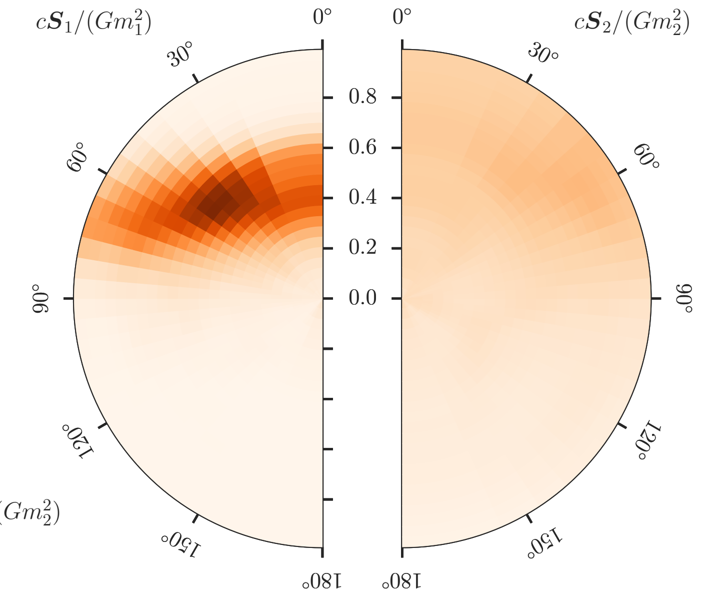
Stellar mass black holes



Spins

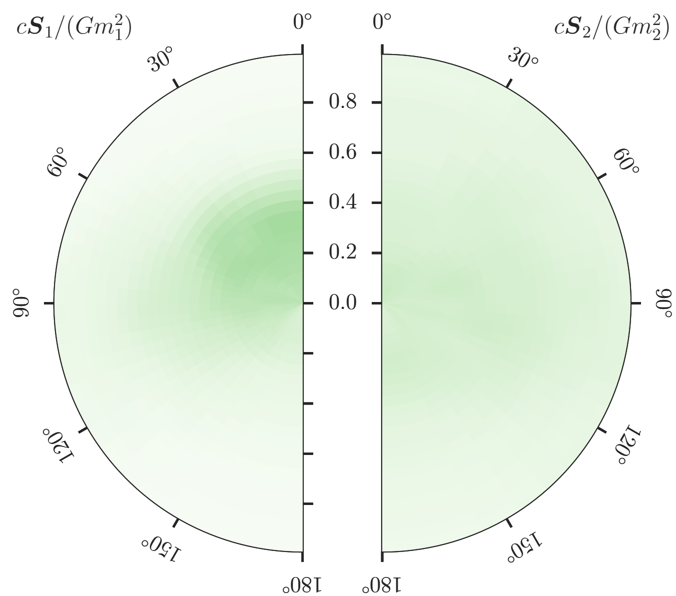


GW150914

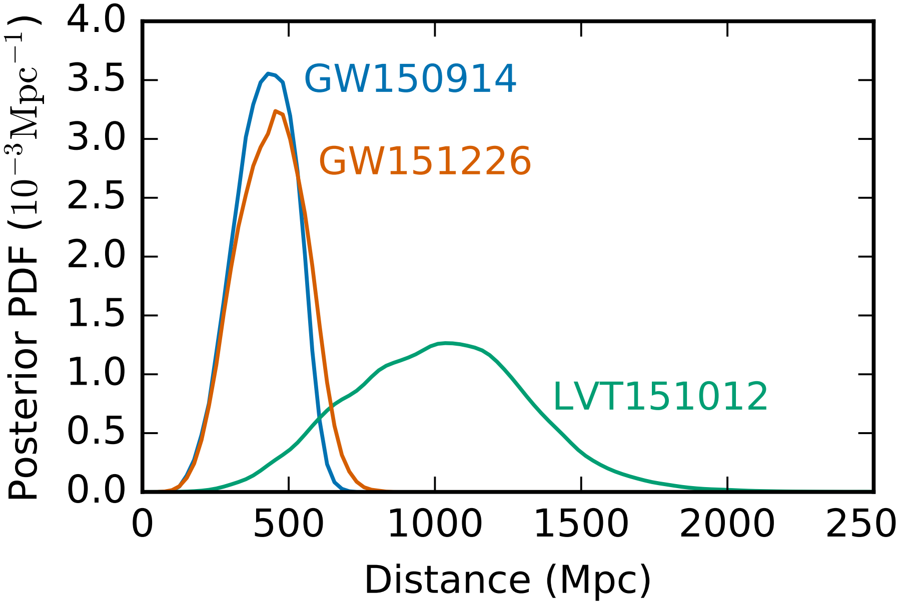


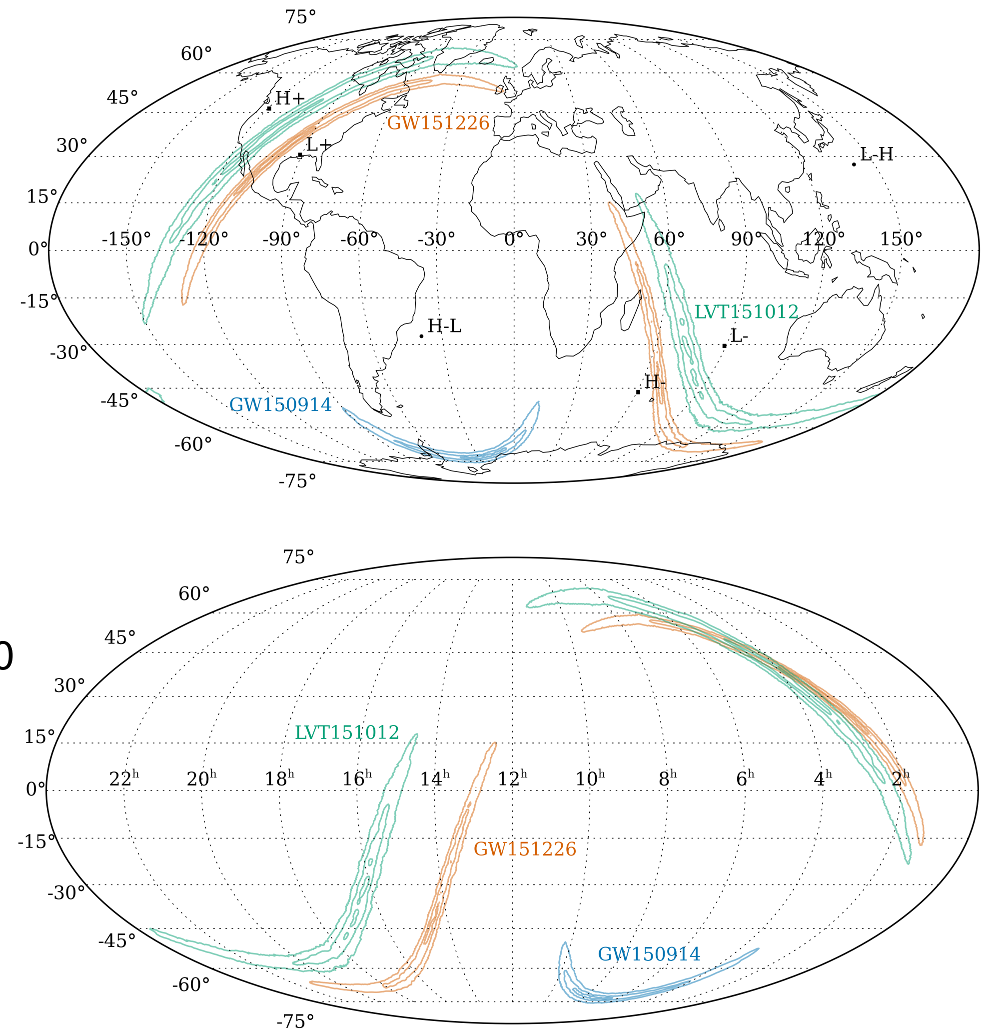
GW151226

LVT151012

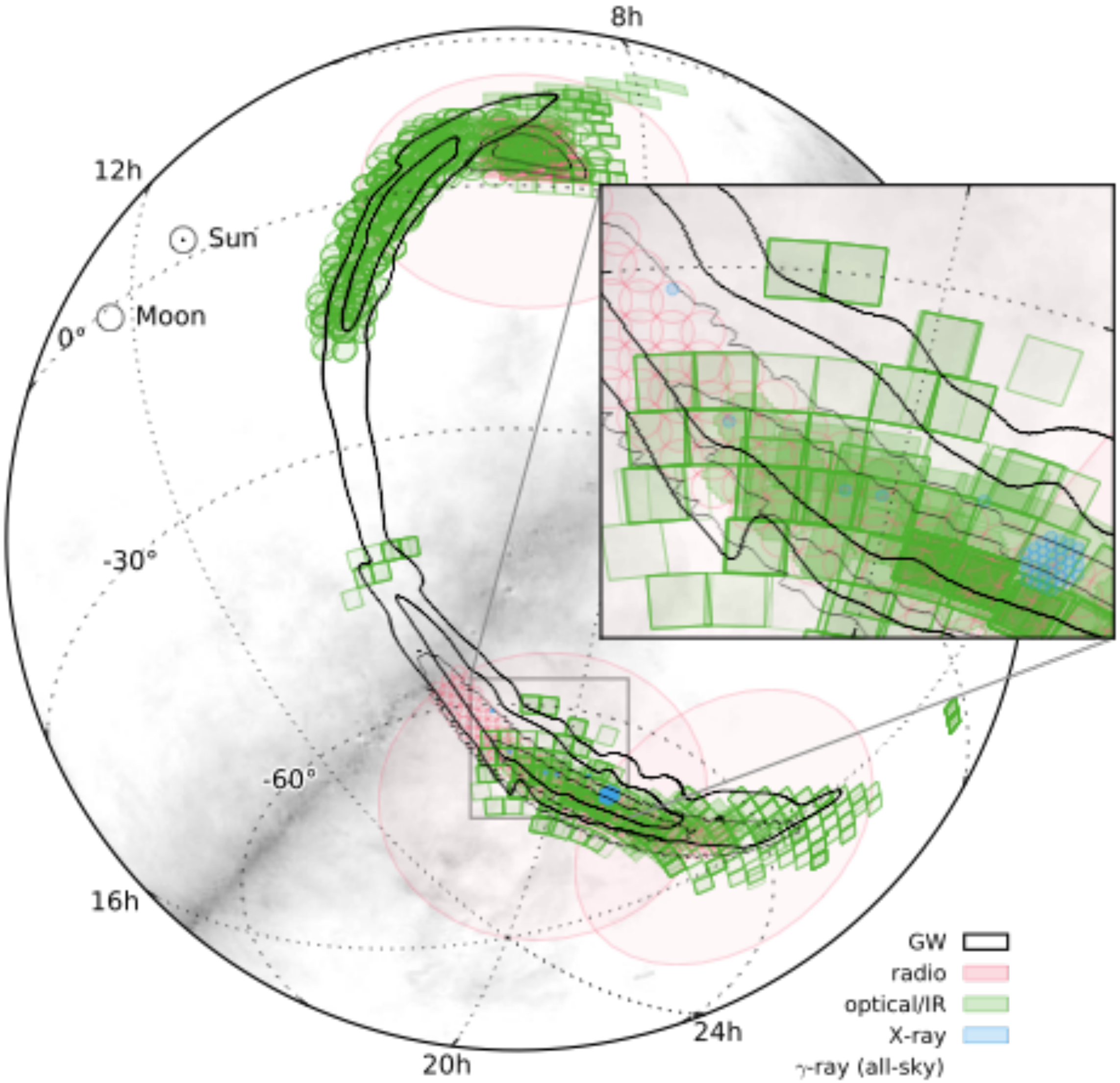


Distance and sky position

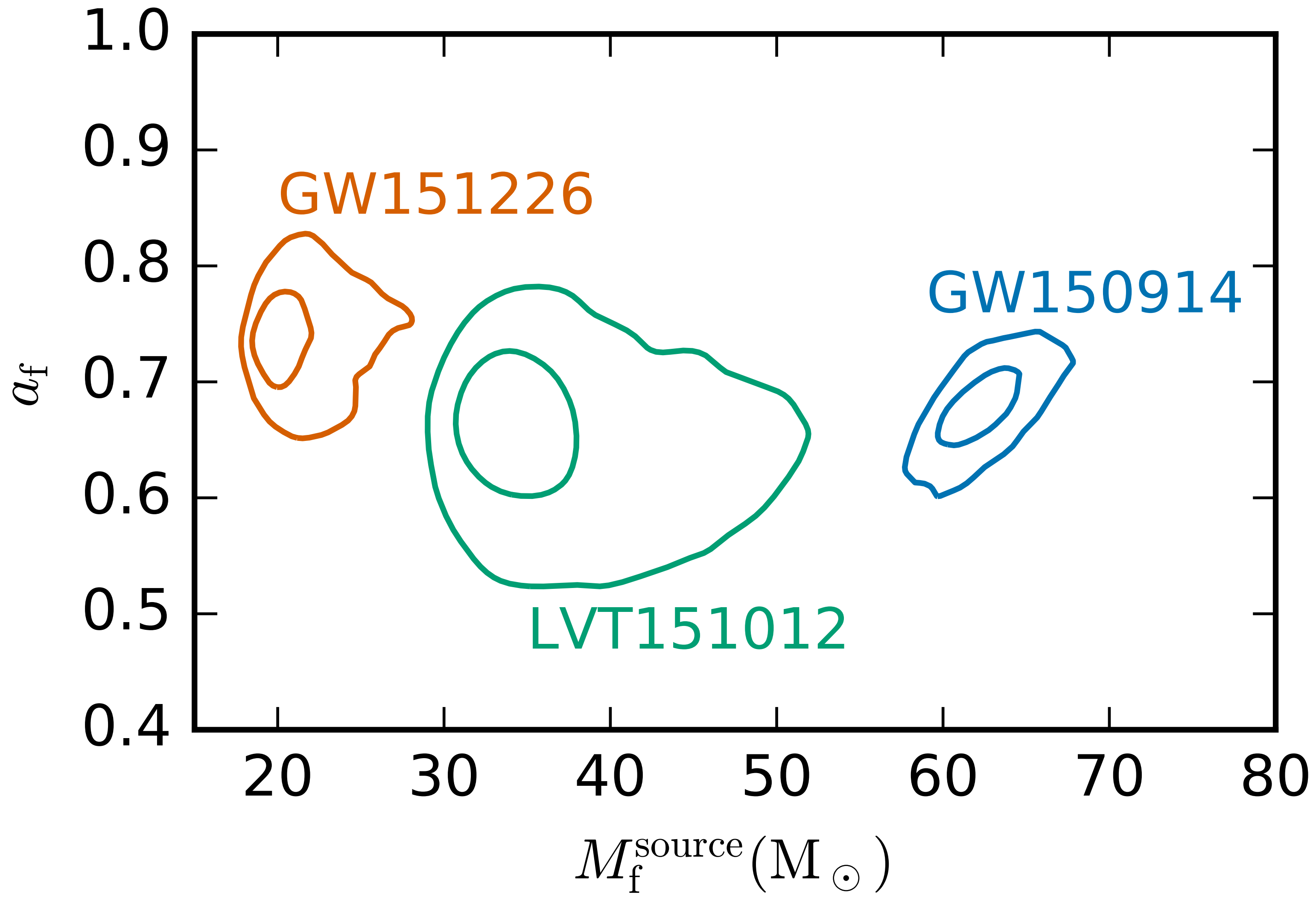




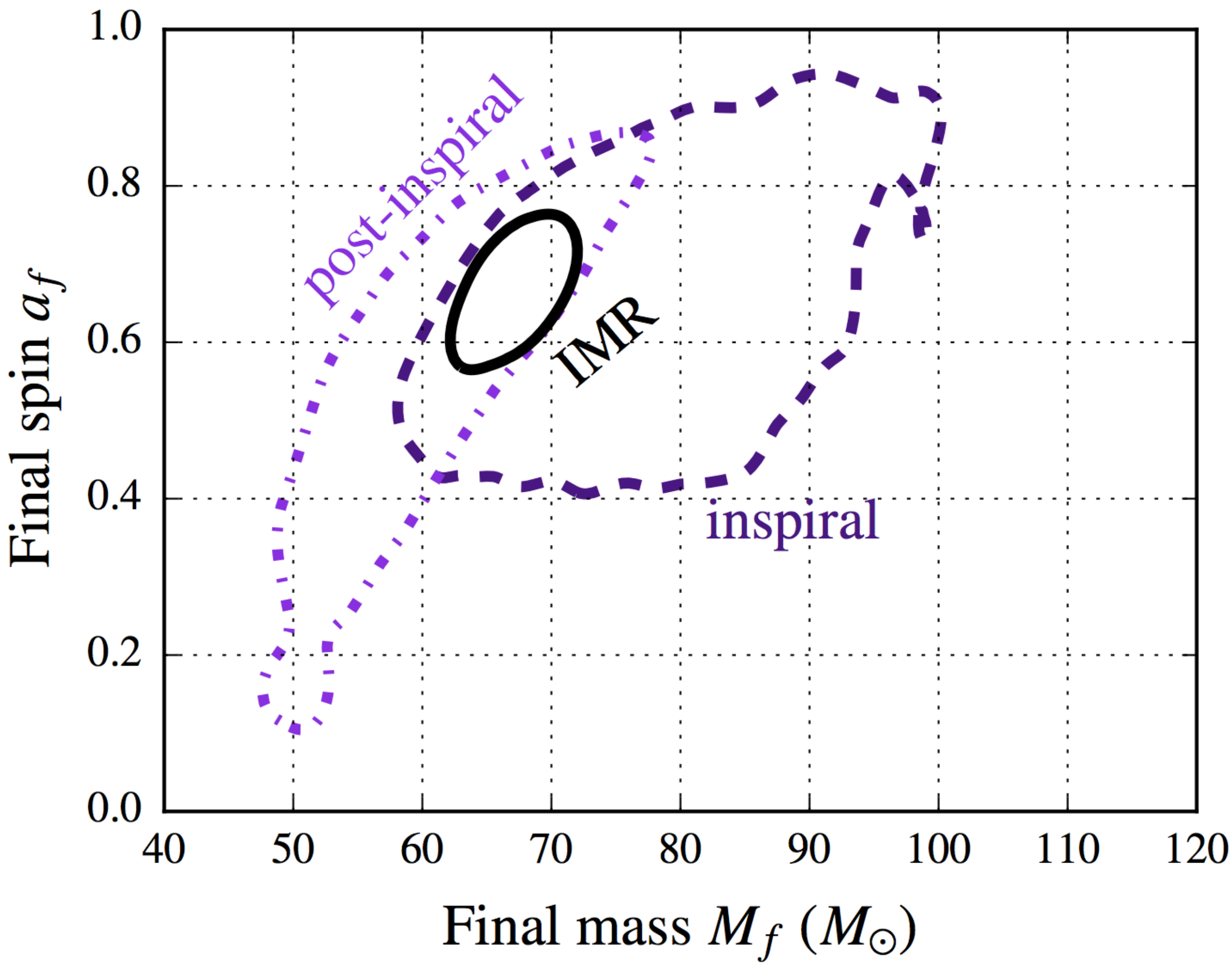
Telescope Observations



Final mass and spin

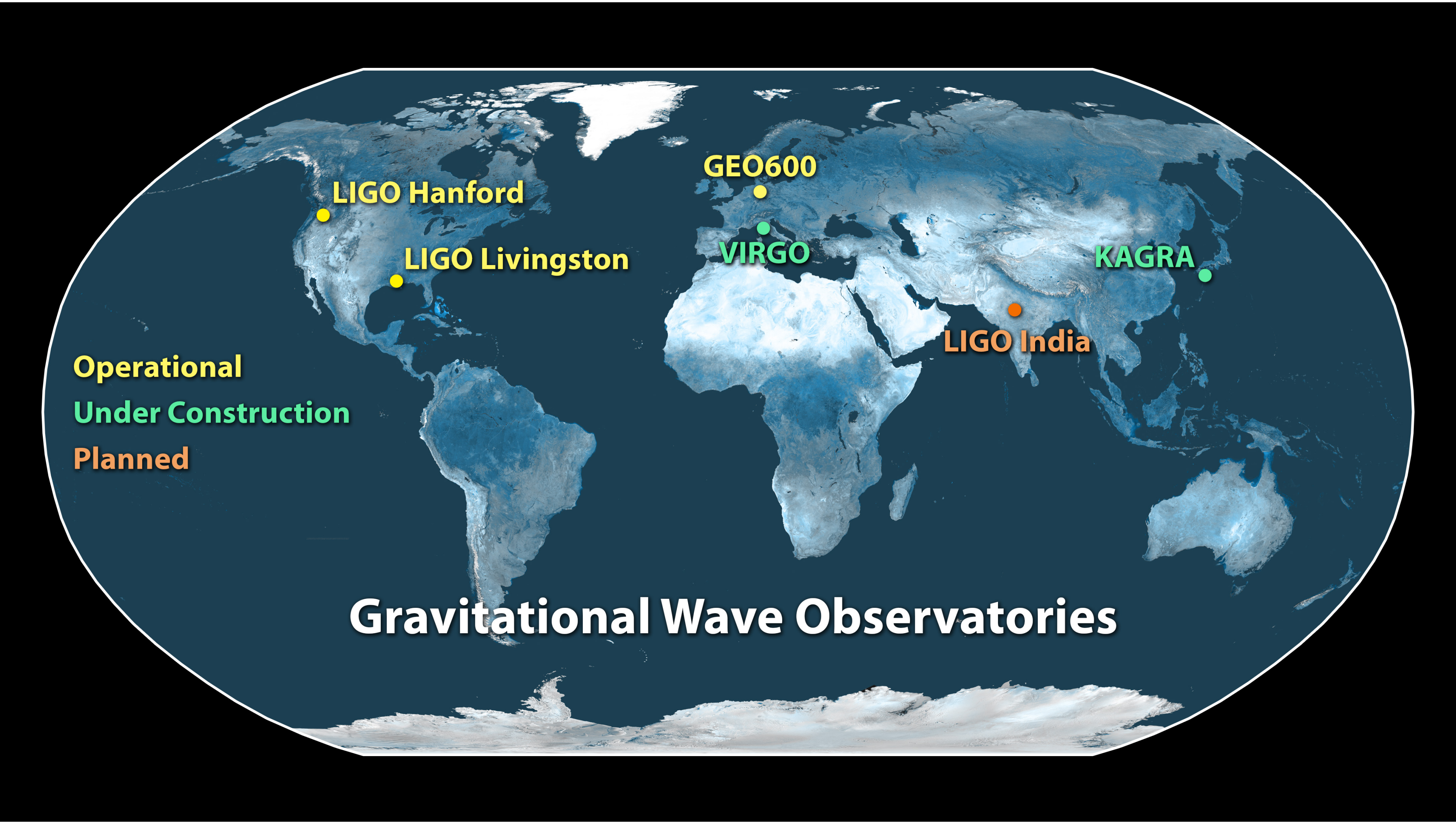


Consistency with General Relativity

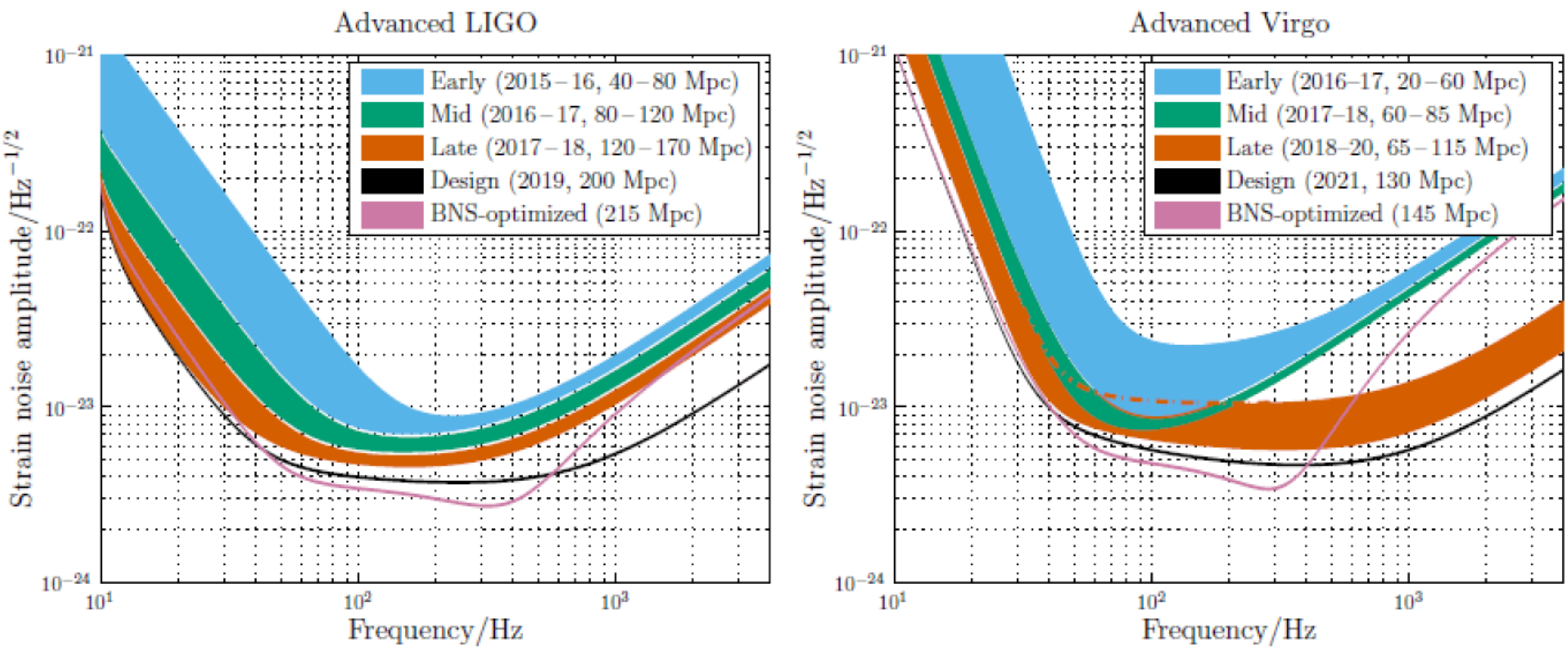


From Abbott et al, "Tests of general relativity with GW150914", 2016

Future Observing

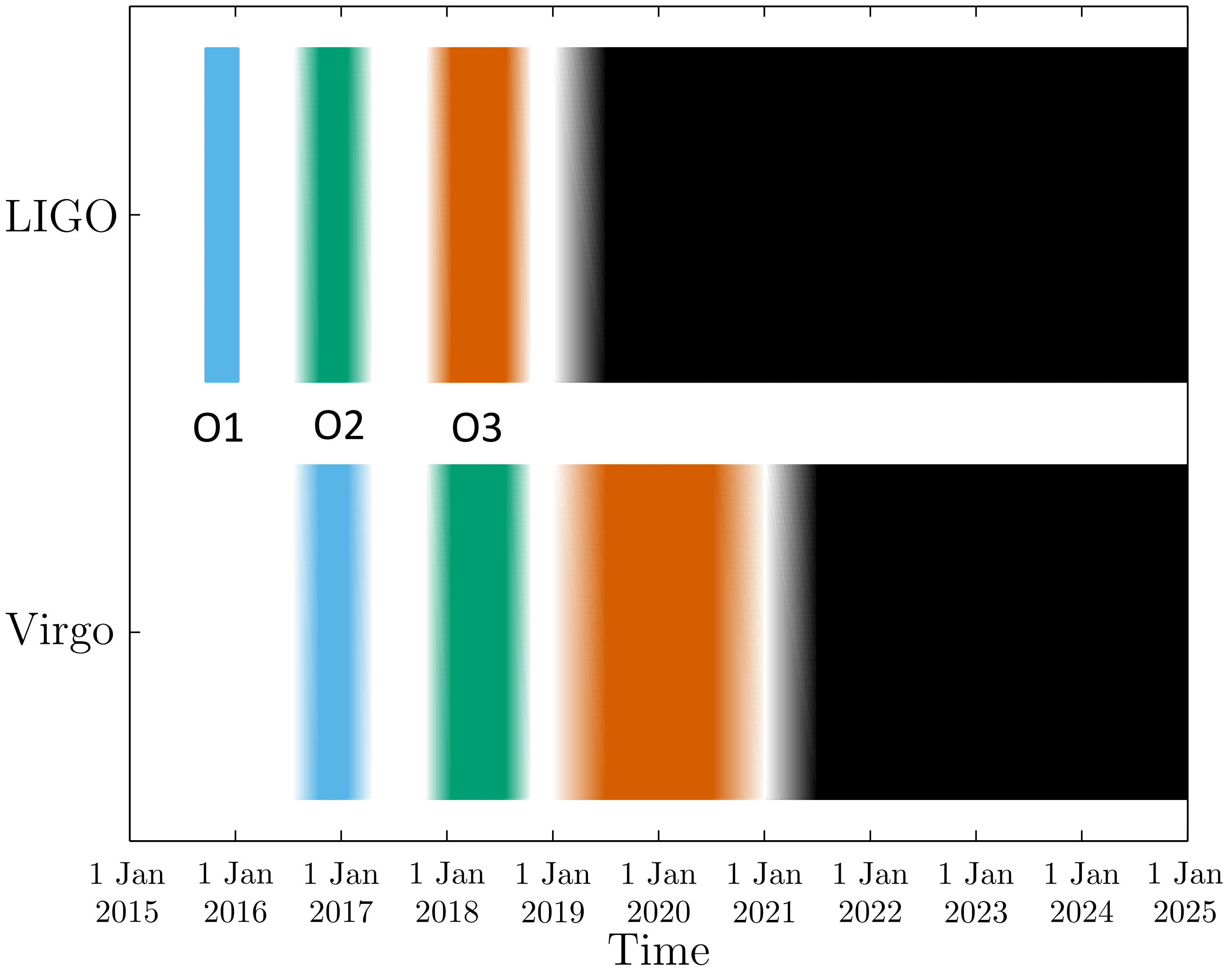


Planned LIGO-Virgo Observing



[Living Rev. Relativity 19 (2016), 1]

Planned LIGO-Virgo Observing



orientation-averaged distance for BNS detection with SNR = 8

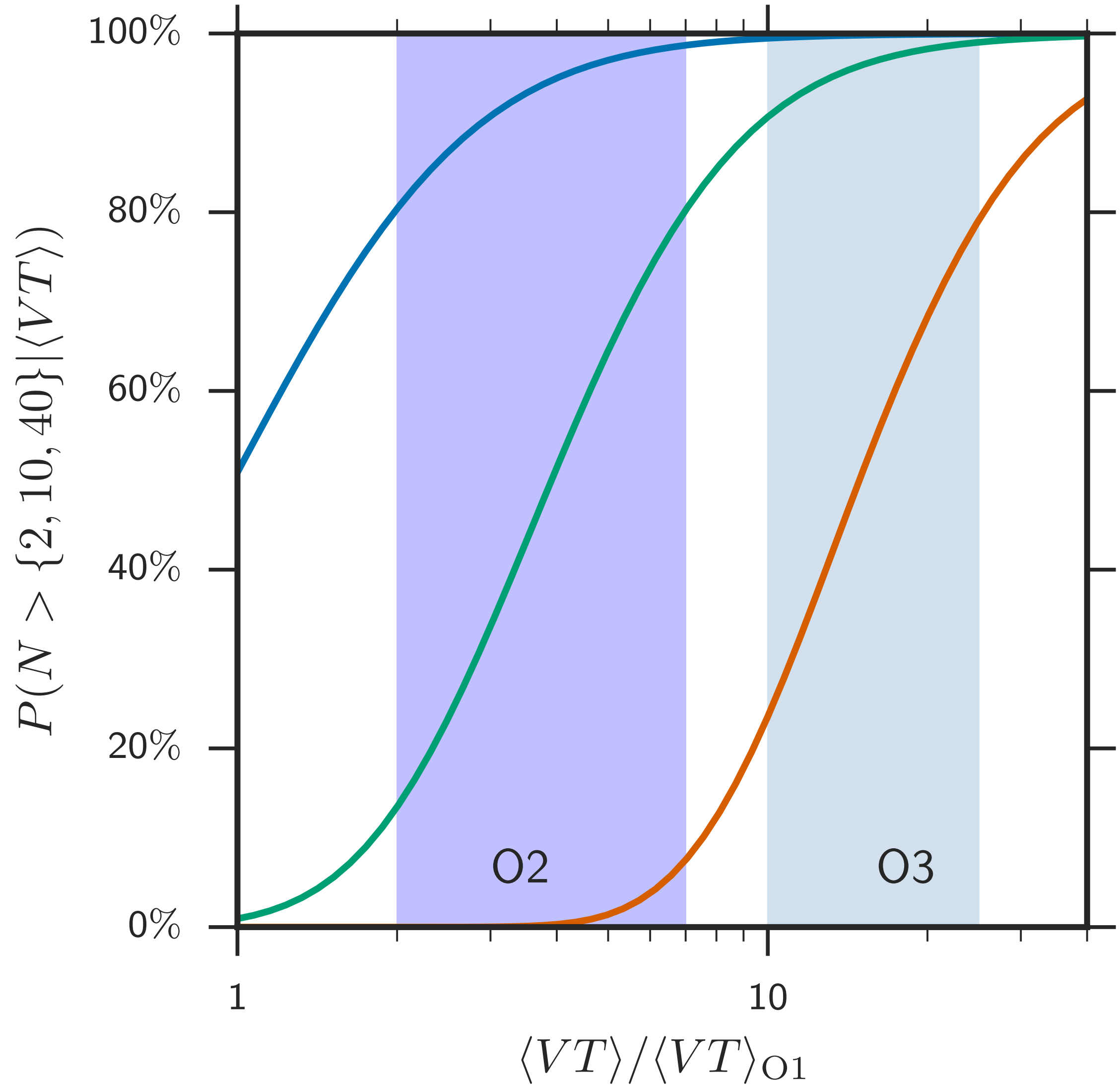
Living Rev. Relativity 19 (2016), 1

Expectations for future runs

Probability of observing

- N > 2 (blue)
- N > 10 (green)
- N > 40 (red)

highly significant events, as a function of surveyed time-volume.



Summary

- GW150914 and GW151226 are the *first direct detections* of GWs and the *first observations of binary black hole mergers*.
- GW150914 contains the *most massive known stellar-mass black holes*.
- GW150914 and GW151226 provide the opportunity *test General Relativity* in the large velocity, highly nonlinear regime.
- LIGO resumed the search for gravitational waves on November 30, 2016.
- We expect to observe many more binary black hole mergers in the coming years, as well as binaries containing neutron stars.
- Continue to look for electromagnetic counterparts to gravitational wave signals.